Binary neutron stars: from macroscopic collisions to microphysics

Luciano Rezzolla

Institute for Theoretical Physics, Frankfurt

Tucson [Madrid] 15.06.22

Theoretical Physics Colloquium

hosted by Prof. Igor Shovkovy at the Arizona State University

Plan of the talk

- The richness of merging binary neutron stars
- GW spectroscopy: EOS from frequencies
- GW170817: a game changer
- Signatures of quark-hadron phase transitions
- On the sound speed in neutron stars
- Threshold mass to prompt collapse
- EM counterparts, ejecta, and jets

The two-body problem in GR

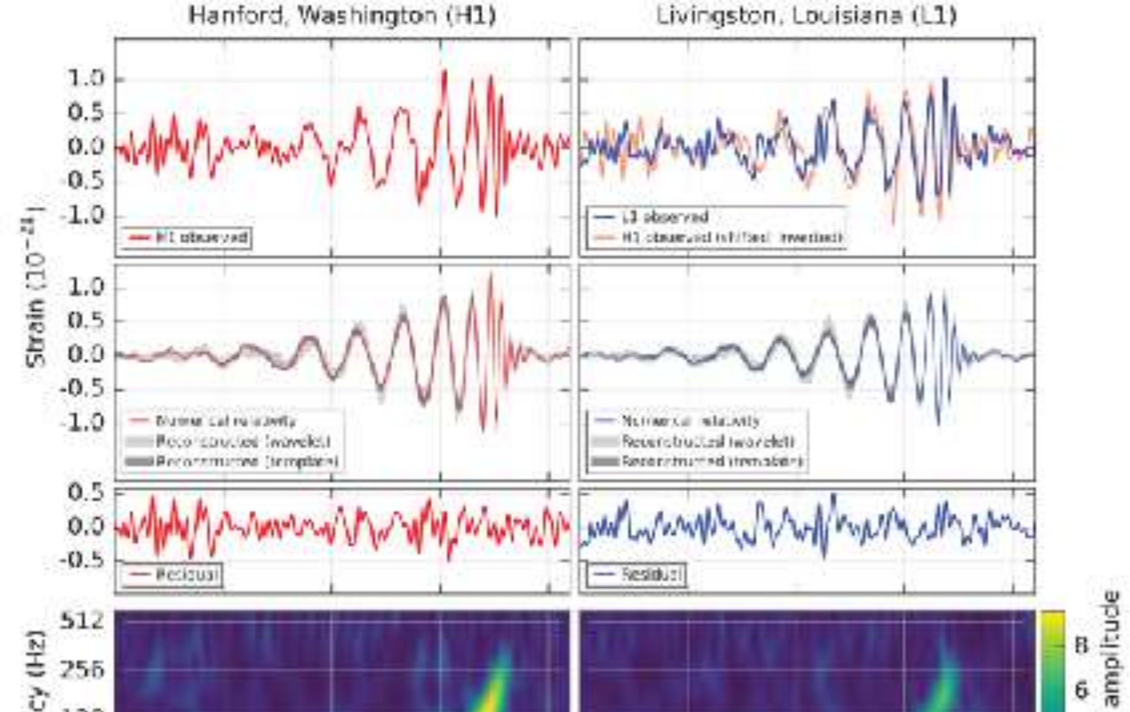
- For black holes the process is very **simple:**

BH + BH → BH + GWs

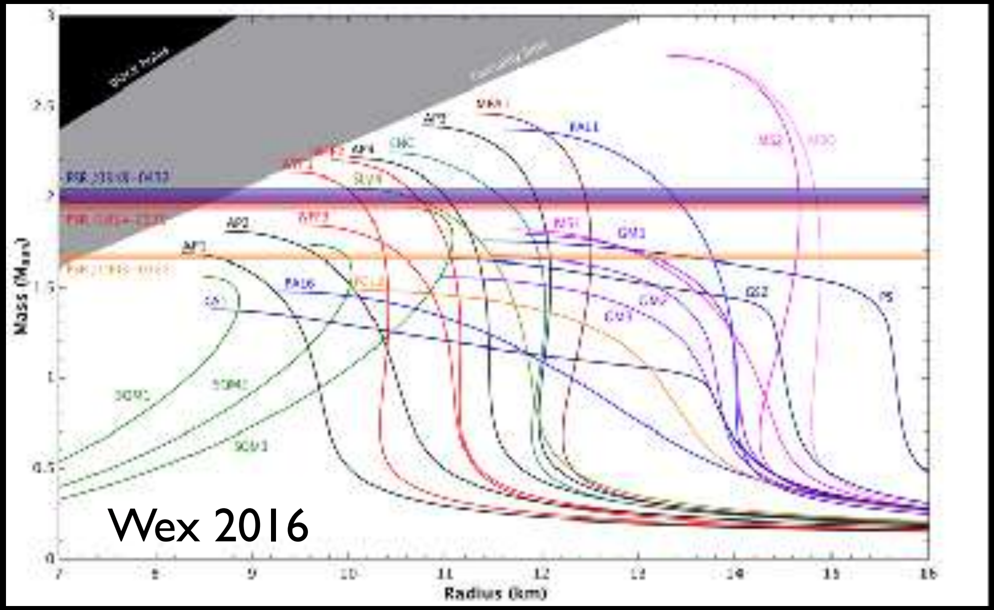
- For NSs the question is more **subtle:** hyper-massive neutron star (HMNS), ie

NS + NS → HMNS+... ? → BH+to

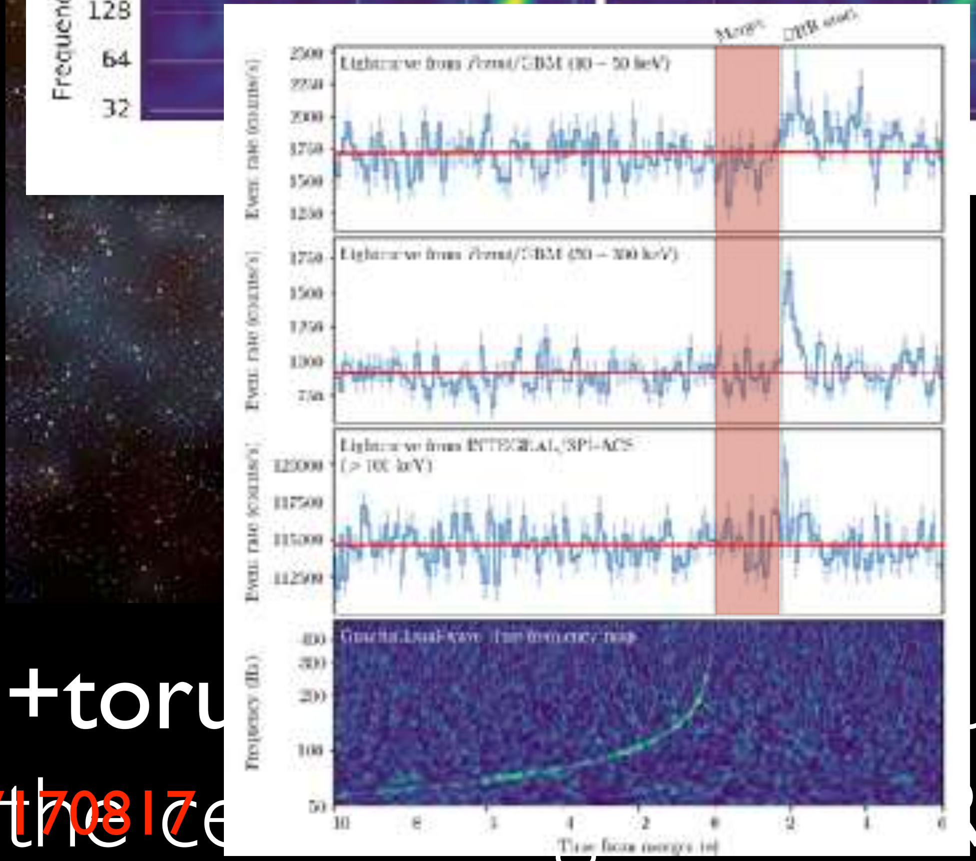
- **HMNS** phase can provide clear information on **EOS**



GW150914



Wex 2016



- BH+toru us on the ce RBs

GW170817

The two-body problem in GR

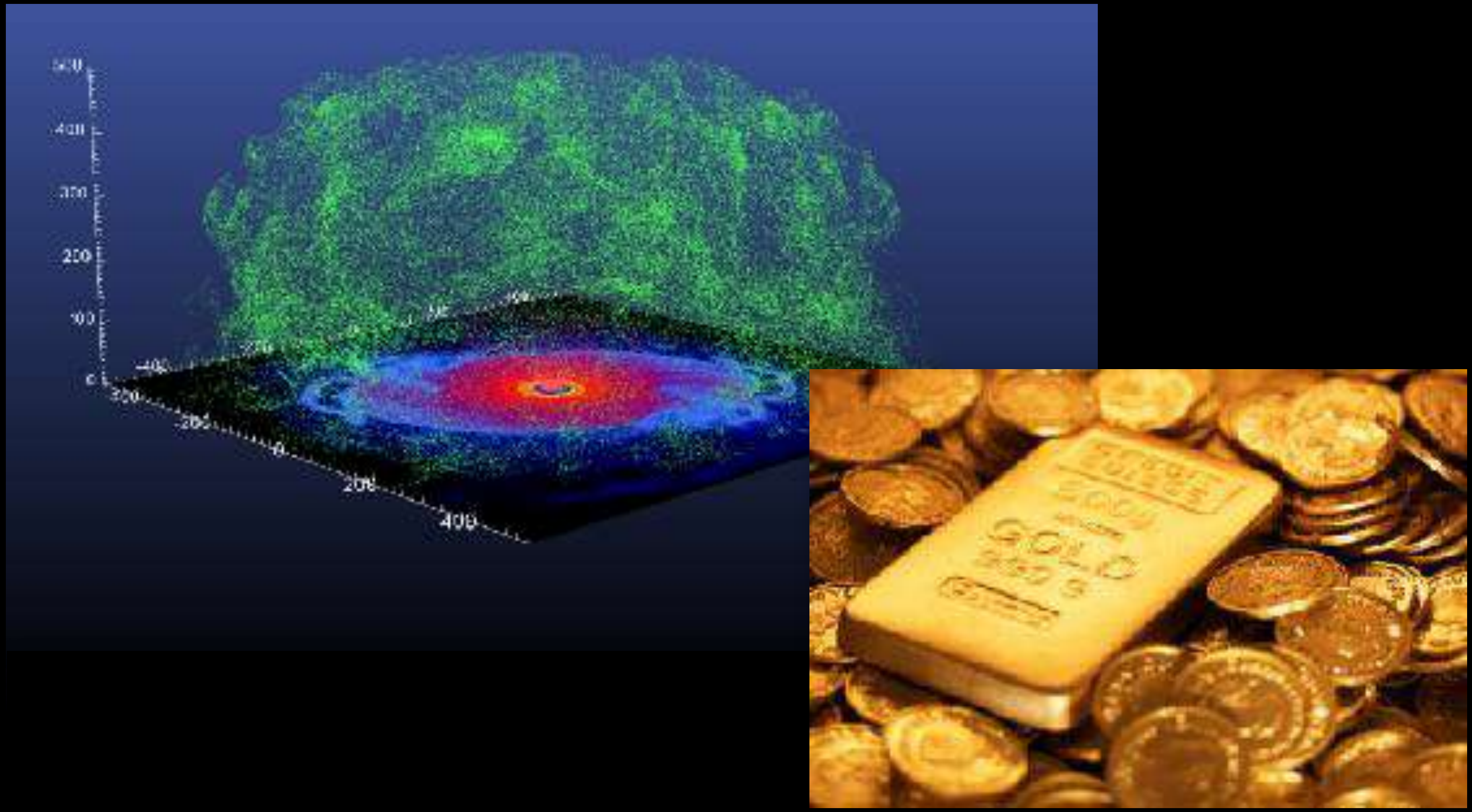
- For black holes the process is very **simple**:

BH + BH → BH + GWs

- For NSs the question is more **subtle:** the merger leads to an hyper-massive neutron star (HMNS), ie a metastable equilibrium:

NS + NS → HMNS+... ? → BH+torus+... ? → BH + GWs

- **ejected matter** undergoes nucleosynthesis of heavy elements



The equations of numerical relativity

$$R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R = 8\pi T_{\mu\nu} \,, \text{ (Einstein equations)}$$

$$\nabla_\mu T^{\mu\nu} = 0 \,, \quad \text{(cons. energy/momentum)}$$

$$\nabla_\mu (\rho u^\mu) = 0 \,, \quad \text{(cons. rest mass)}$$

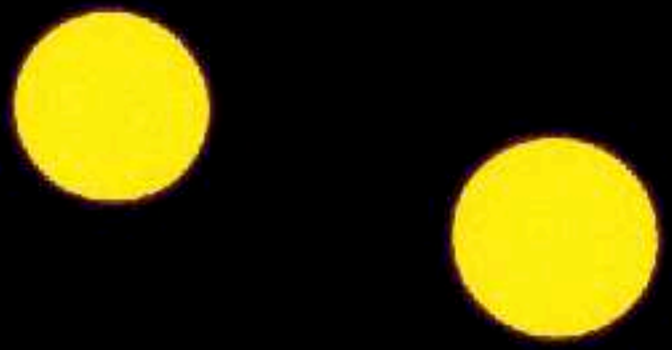
$$p = p(\rho, \epsilon, Y_e, \ldots) \,, \quad \text{(equation of state)}$$

$$\nabla_\nu F^{\mu\nu} = I^\mu \,, \qquad \nabla_\nu^* F^{\mu\nu} = 0 \,, \quad \text{(Maxwell equations)}$$

$$T_{\mu\nu} = T_{\mu\nu}^{\text{fluid}} + T_{\mu\nu}^{\text{EM}} + \ldots \quad \text{(energy} - \text{momentum tensor)}$$

Animations: Breu, Radice, LR

A prototypical simulation with possibly the best code looks like this…



merger → HMNS → BH + torus

$M = 2 \times 1.35\, M_\odot$

LS220 EOS

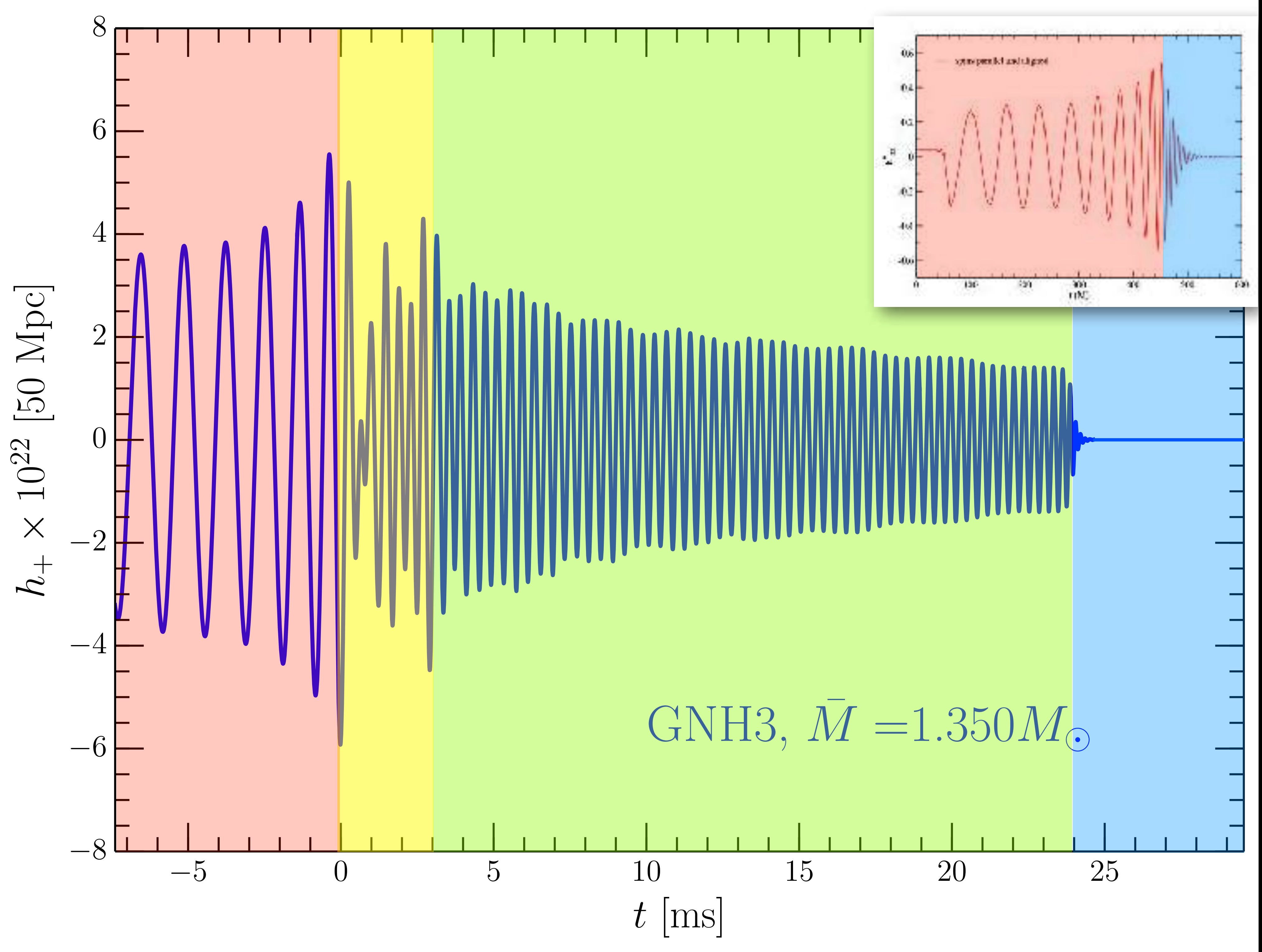
Qualitatively, this is what normally happens:

merger → HMNS → BH + torus

Quantitatively, differences are produced by:

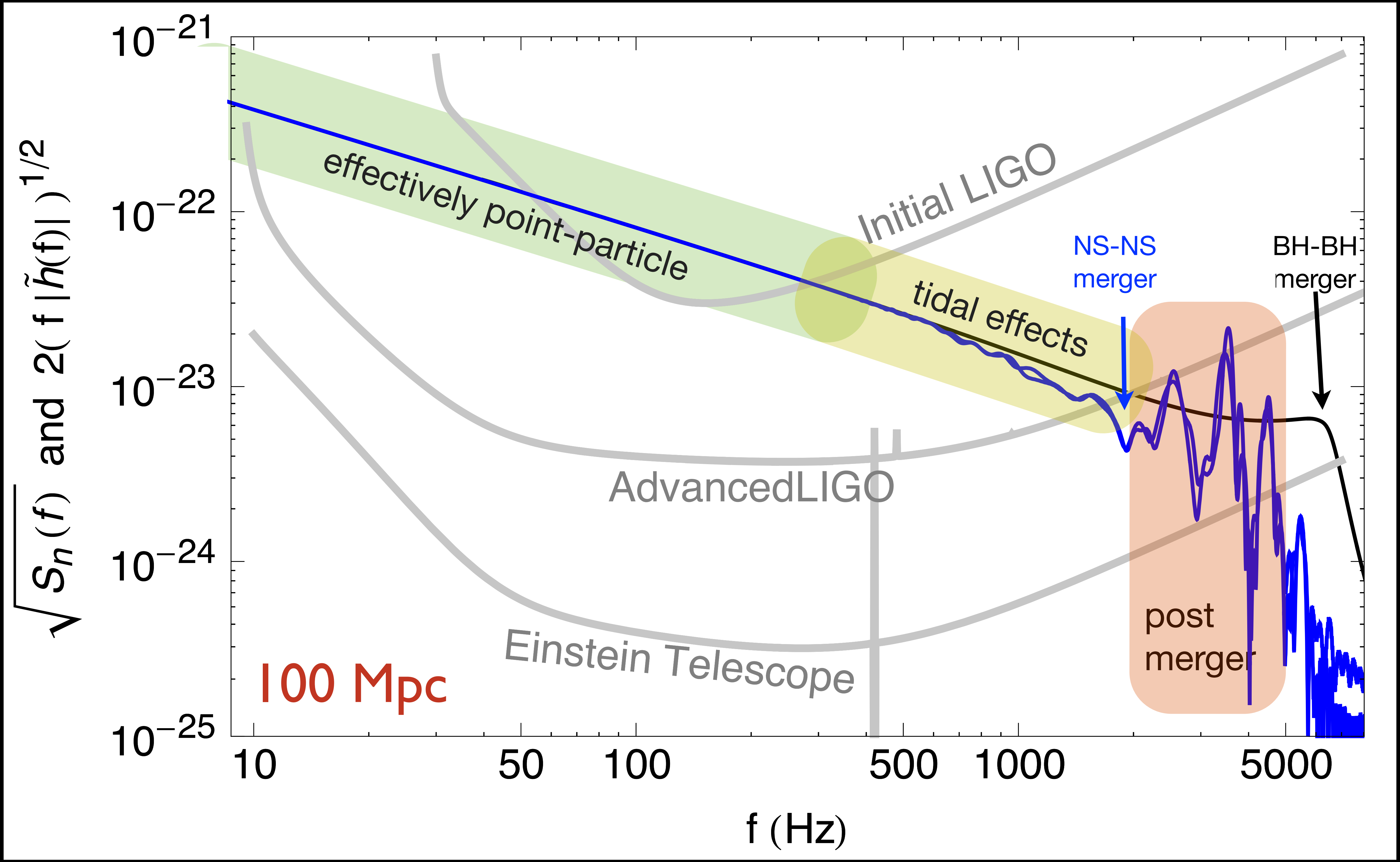
- total mass (prompt vs delayed collapse)
- mass asymmetries (HMNS and torus)
- soft/stiff EOS (inspiral and post-merger, PT)
- magnetic fields (equil. and EM emission)
- radiative losses (equil. and nucleosynthesis)

Anatomy of the GW signal



Postmerger signal: peculiar of binary NSs

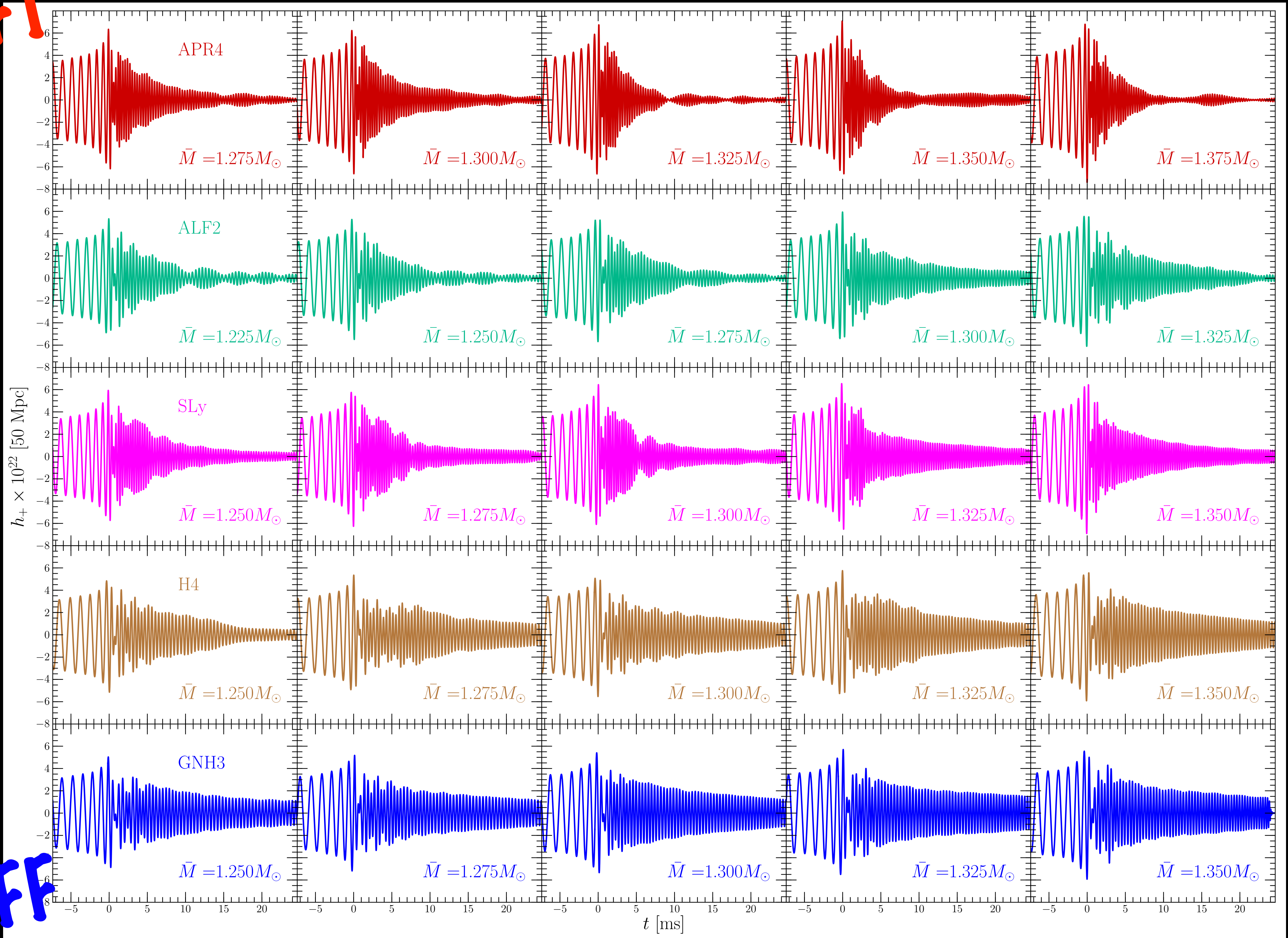
In frequency space



Read et al. (2013)

What we can do nowadays

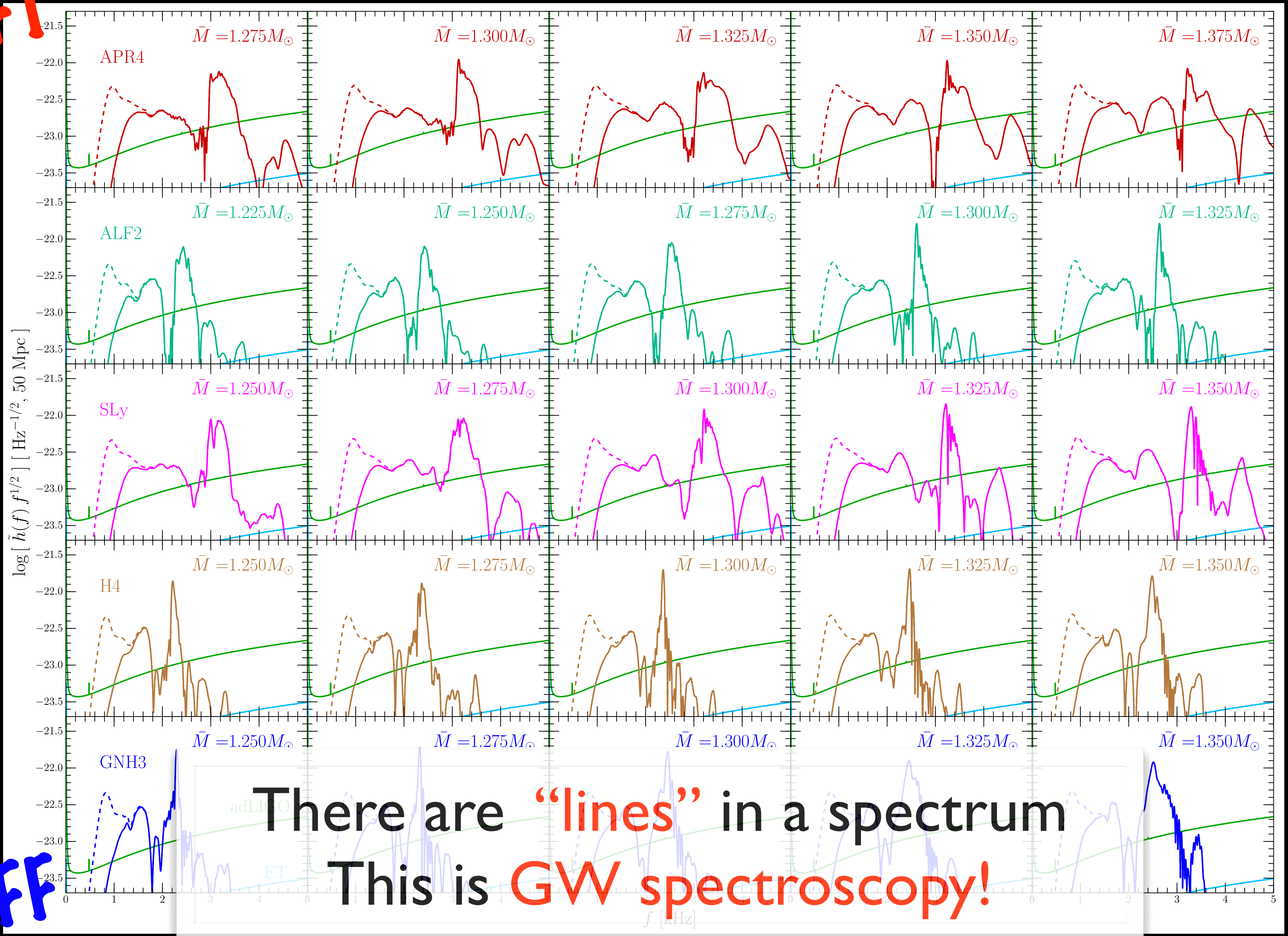
Takami, LR, Baiotti (2014, 2015), LR+ (2016)



Extracting information from the EOS

Takami, LR, Baiotti (2014, 2015), LR+ (2016)

SOFT



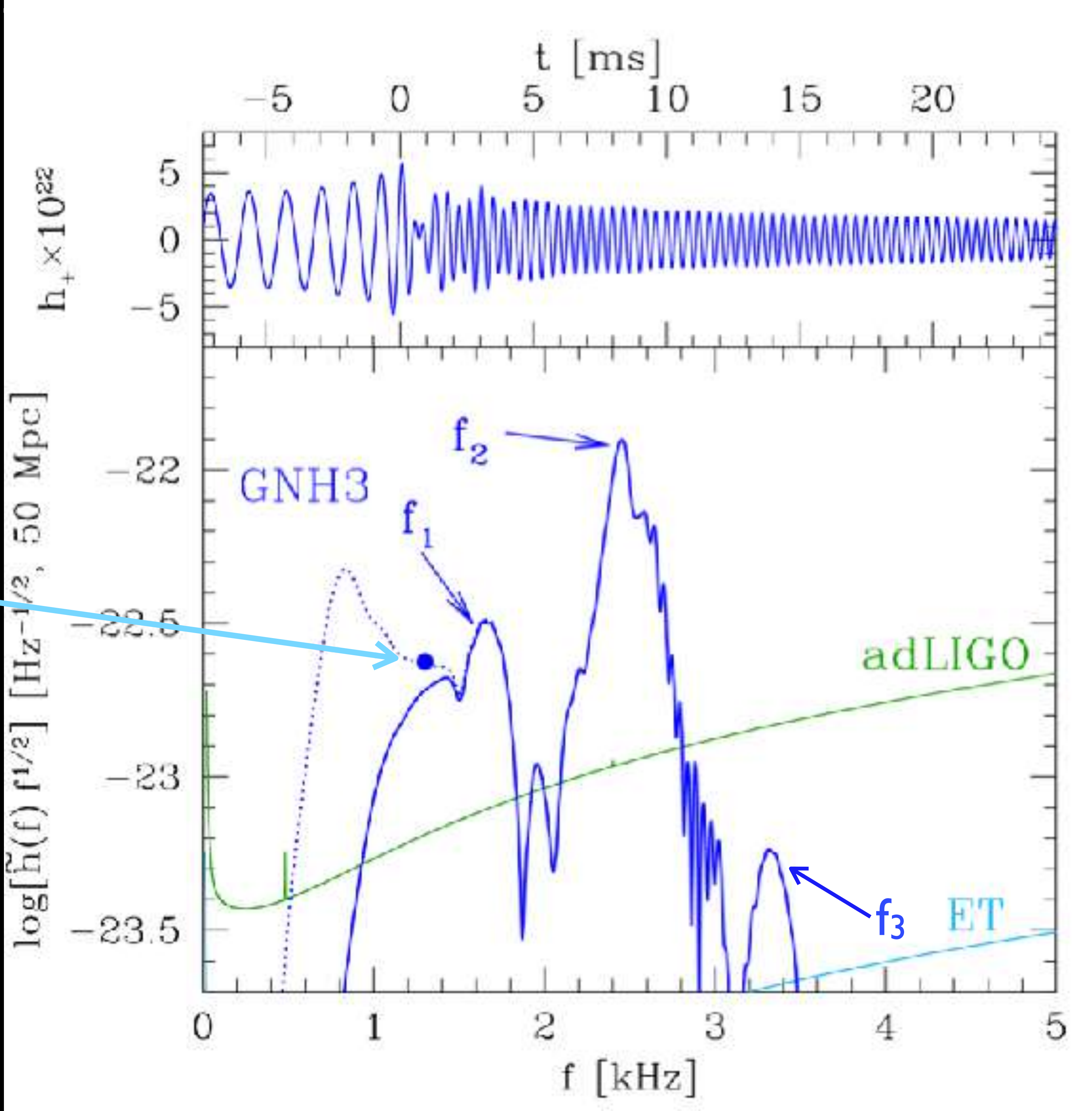
STIFF

There are "lines" in a spectrum

This is GW spectroscopy!

A spectroscopic approach to the EOS

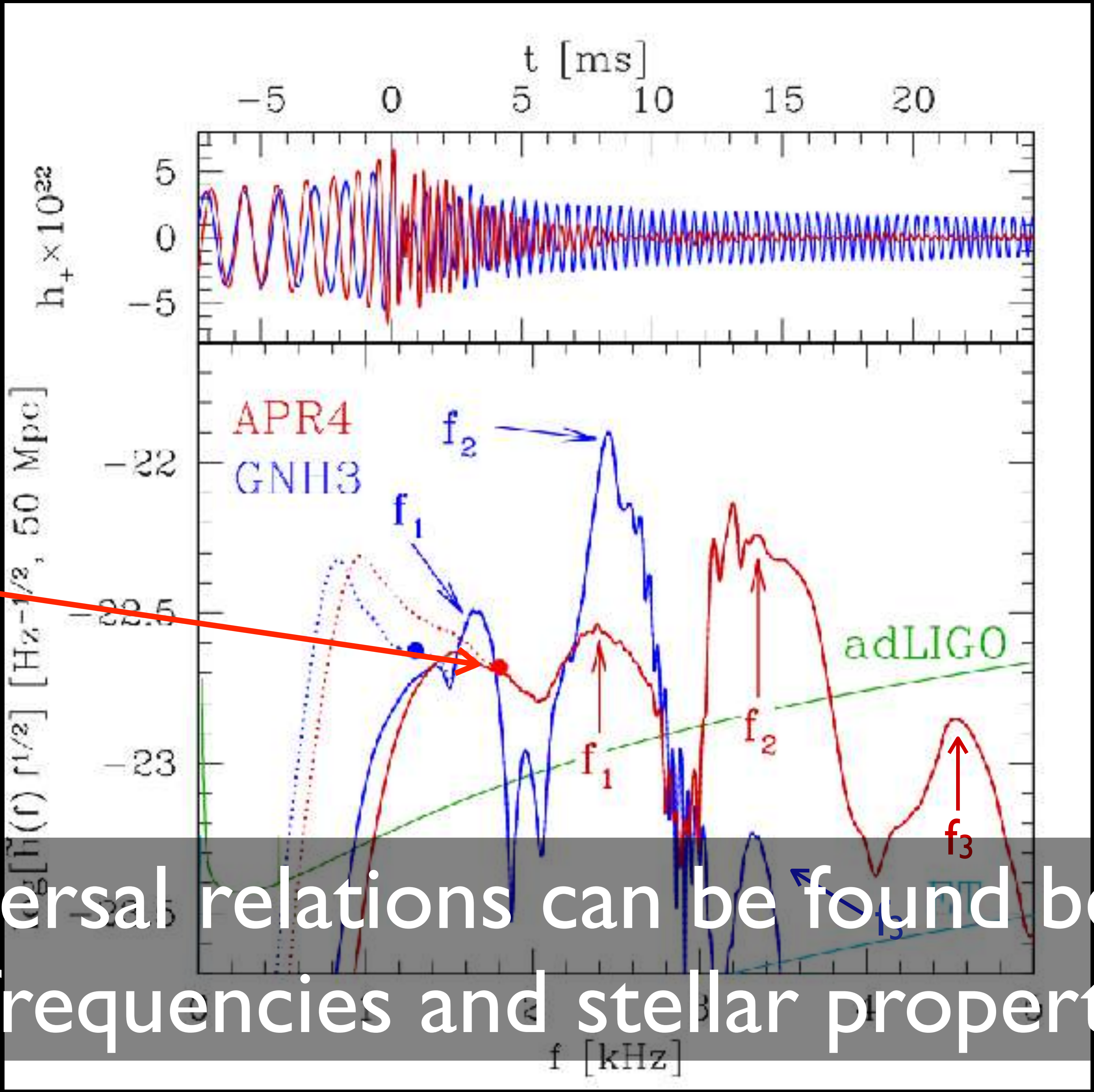
Oechslin+2007, Baiotti+2008, Bauswein+ 2011, 2012, Stergioulas+ 2011, Hotokezaka+ 2013, Takami 2014, 2015, Bernuzzi 2014, 2015, Bauswein+ 2015, Clark+ 2016, LR+2016, de Pietri+ 2016, Feo+ 2017, Bose+ 2017 .



merger frequency

A spectroscopic approach to the EOS

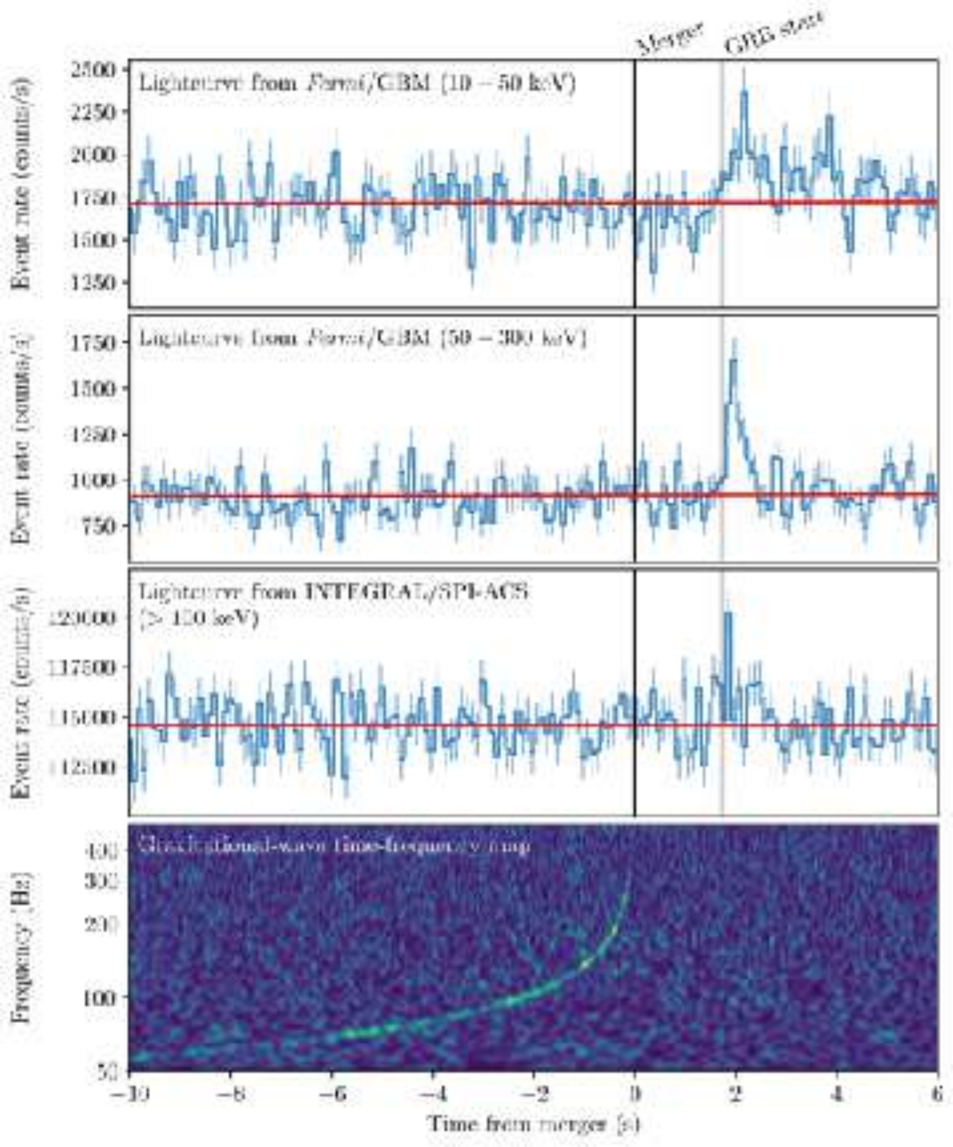
Oechslin+2007, Baiotti+2008, Bauswein+ 2011, 2012, Stergioulas+ 2011, Hotokezaka+ 2013, Takami 2014, 2015, Bernuzzi 2014, 2015, Bauswein+ 2015, Clark+ 2016, LR+2016, de Pietri+ 2016, Feo+ 2017, Bose+ 2017 .



merger frequency

Universal relations can be found between frequencies and stellar properties

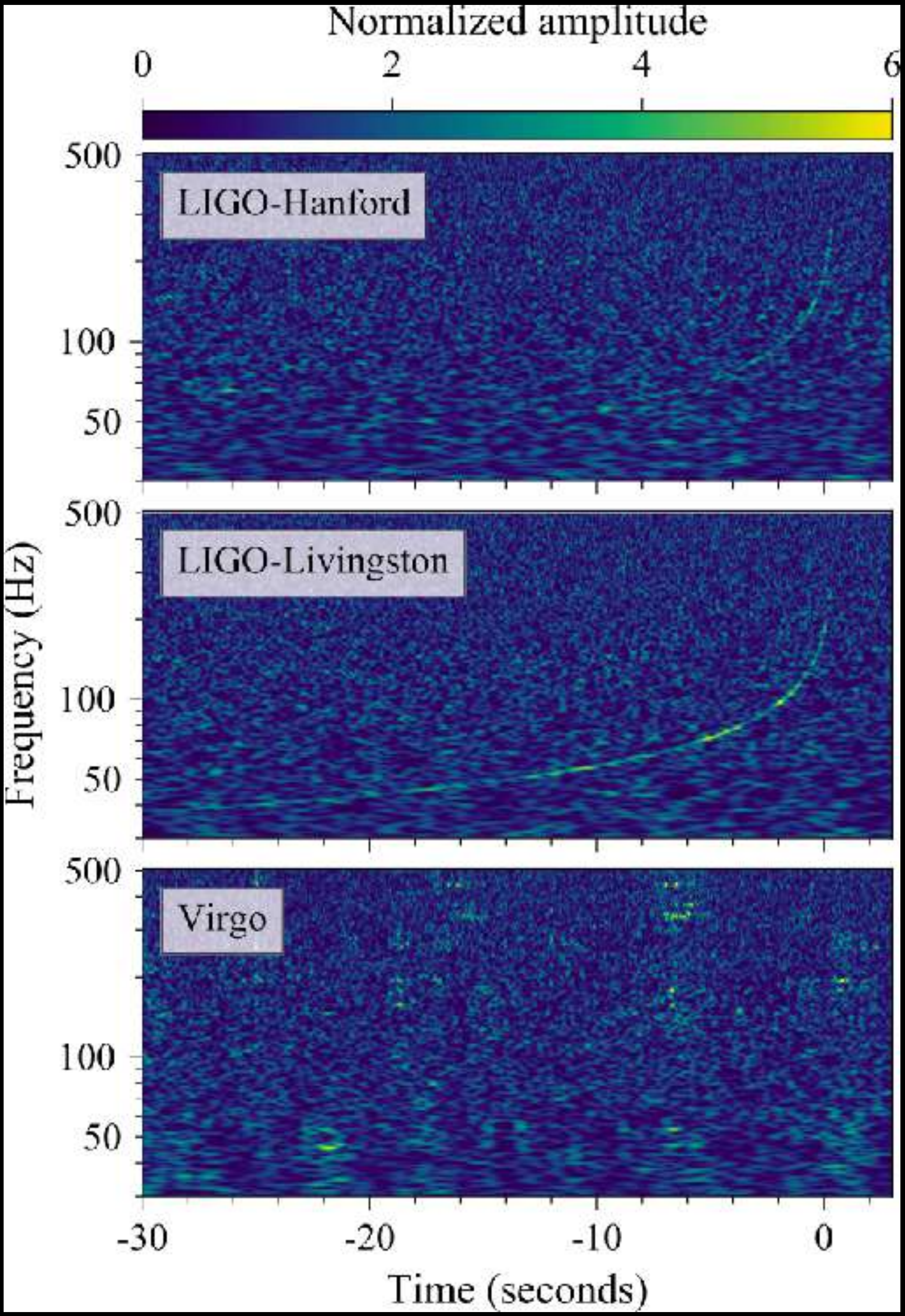
GW170817: a game changer



LR, Most, Weih, ApJL (2018)
Most, Weih, LR, Schaffner-Bielich, PRL (2018)
Nathanail, Most, LR, ApJL (2021)

GW170817: the first binary neutron-star system

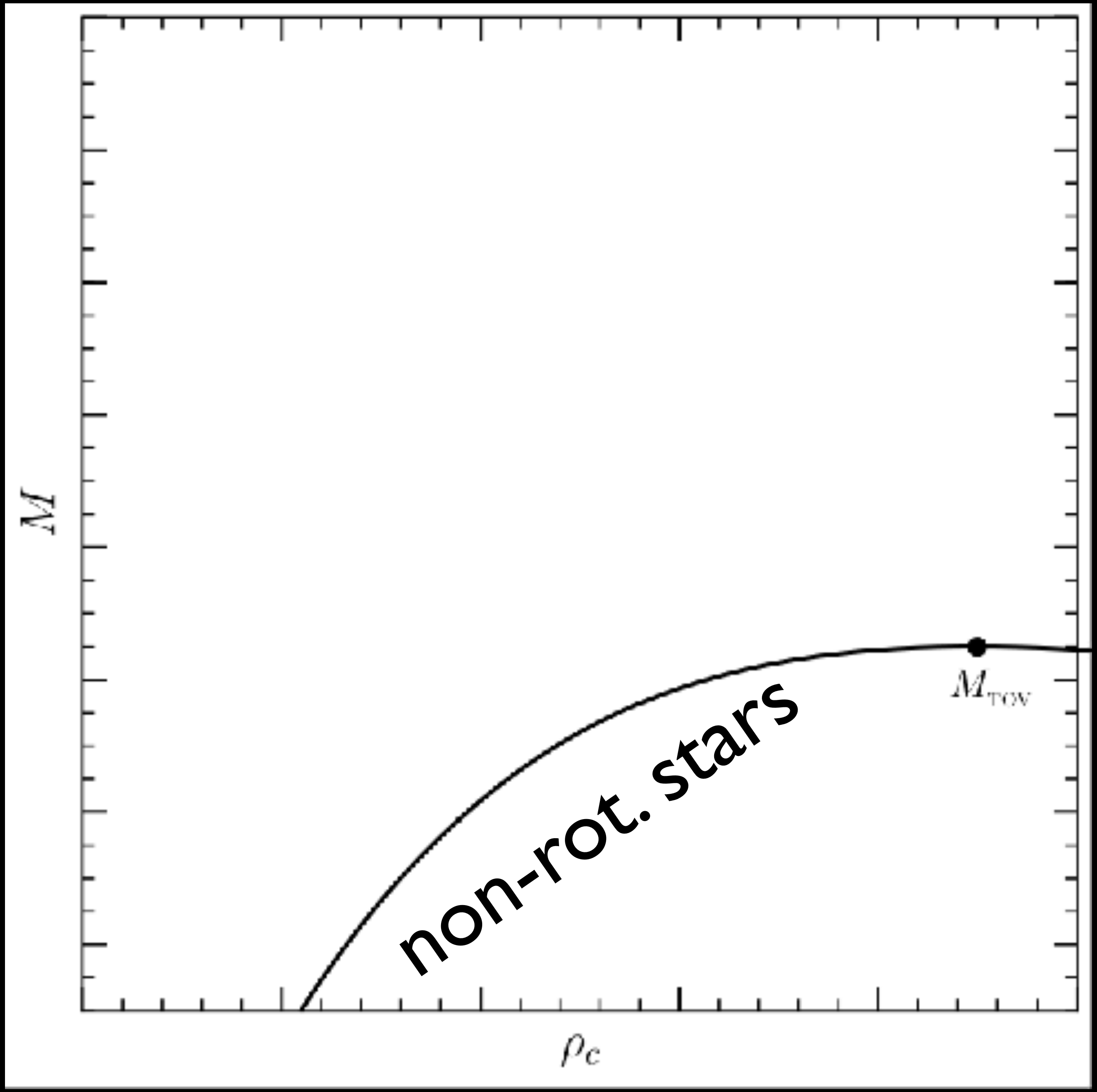
- **Unfortunately** only the **inspiral** signal was detected.
- **Fortunately** this was **sufficient** to set a number of constraints on max. mass, tidal deformability, radii, etc.



Limits on the maximum mass

- The remnant of GW170817 was a hypermassive star, i.e. a differentially rotating object with initial **gravitational** mass:

$$M_1 + M_2 = 2.74^{+0.04}_{-0.01}\, M_\odot$$

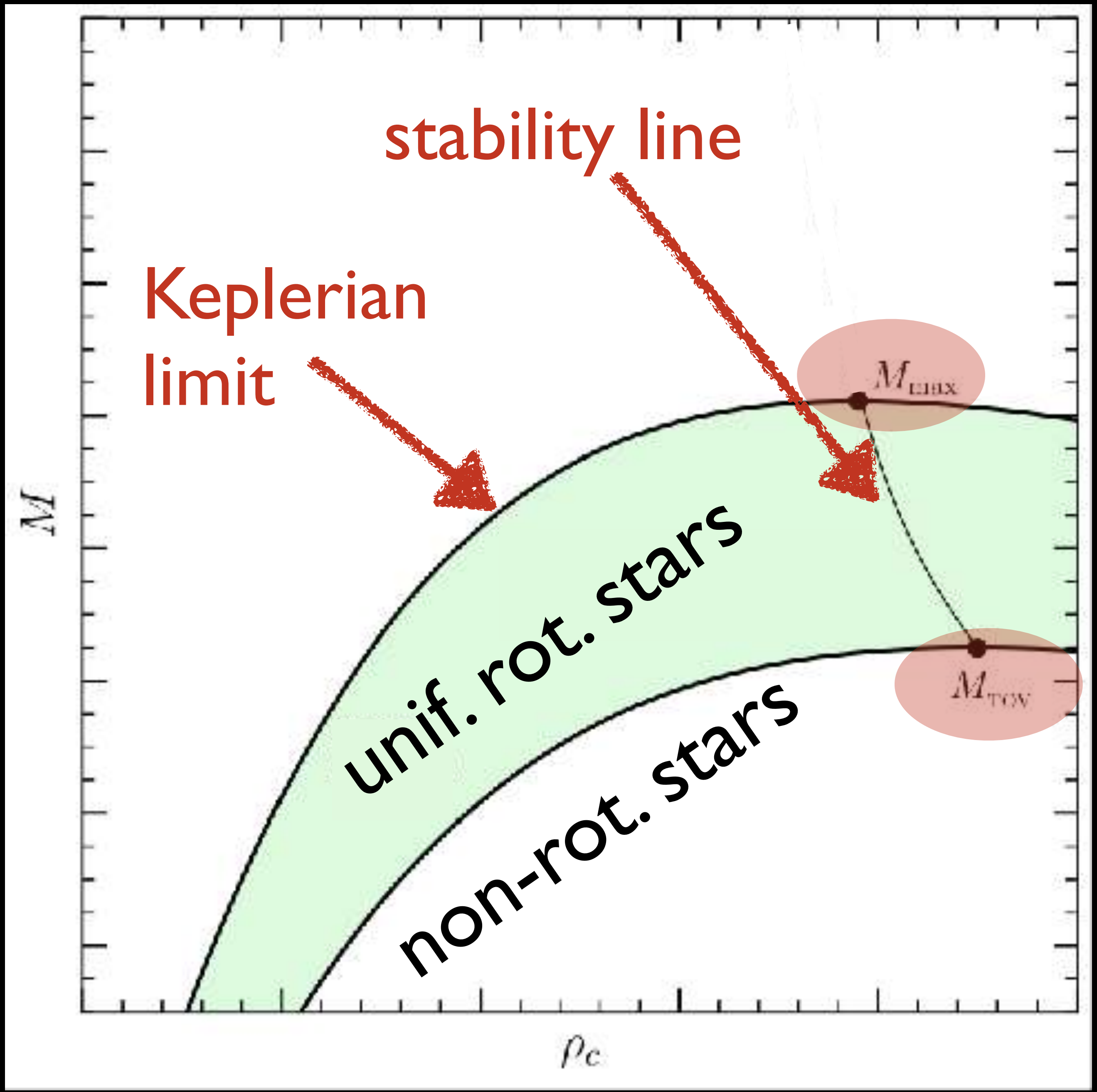


- Sequences of equilibrium models of **nonrotating** stars will have a maximum mass: $M_{\rm TOV}$

Limits on the maximum mass

- The remnant of GW170817 was a hypermassive star, i.e. a differentially rotating object with initial **gravitational** mass:

$$M_1 + M_2 = 2.74^{+0.04}_{-0.01}\, M_\odot$$



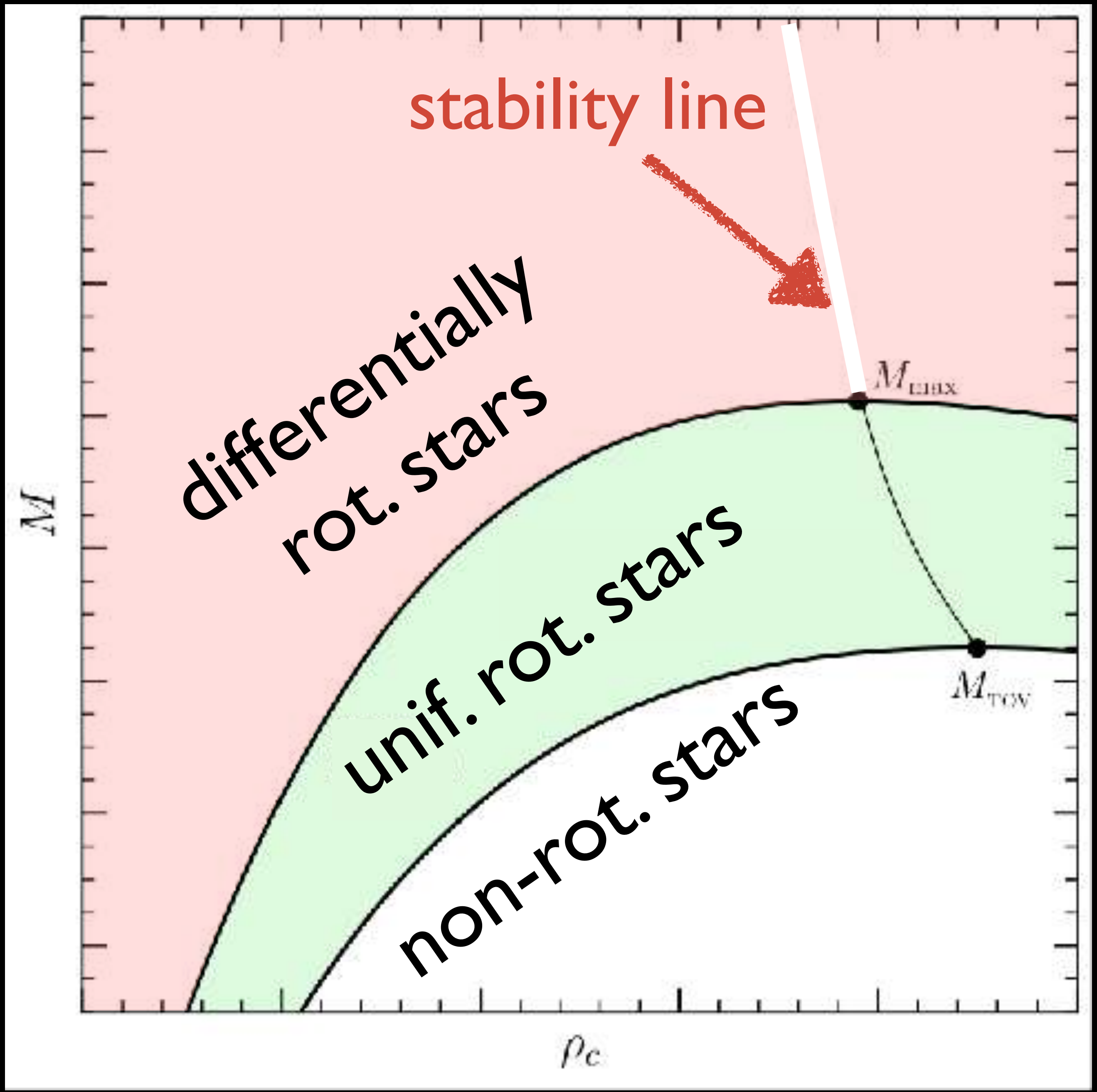
- Sequences of equilibrium models of **nonrotating** stars will have a maximum mass: $M_{\rm TOV}$
- This is true also for **uniformly** rotating stars at mass shedding limit: $M_{\rm max}$
- $M_{\rm max}$ simple and **quasi-universal** function of $M_{\rm TOV}$ (Breu & LR 2016)

$$\boxed{M_{\rm max} = 1.20^{+0.02}_{-0.05}\, M_\odot}$$

Limits on the maximum mass

- The remnant of GW170817 was a hypermassive star, i.e. a differentially rotating object with initial **gravitational** mass:

$$M_1 + M_2 = 2.74^{+0.04}_{-0.01}\, M_\odot$$

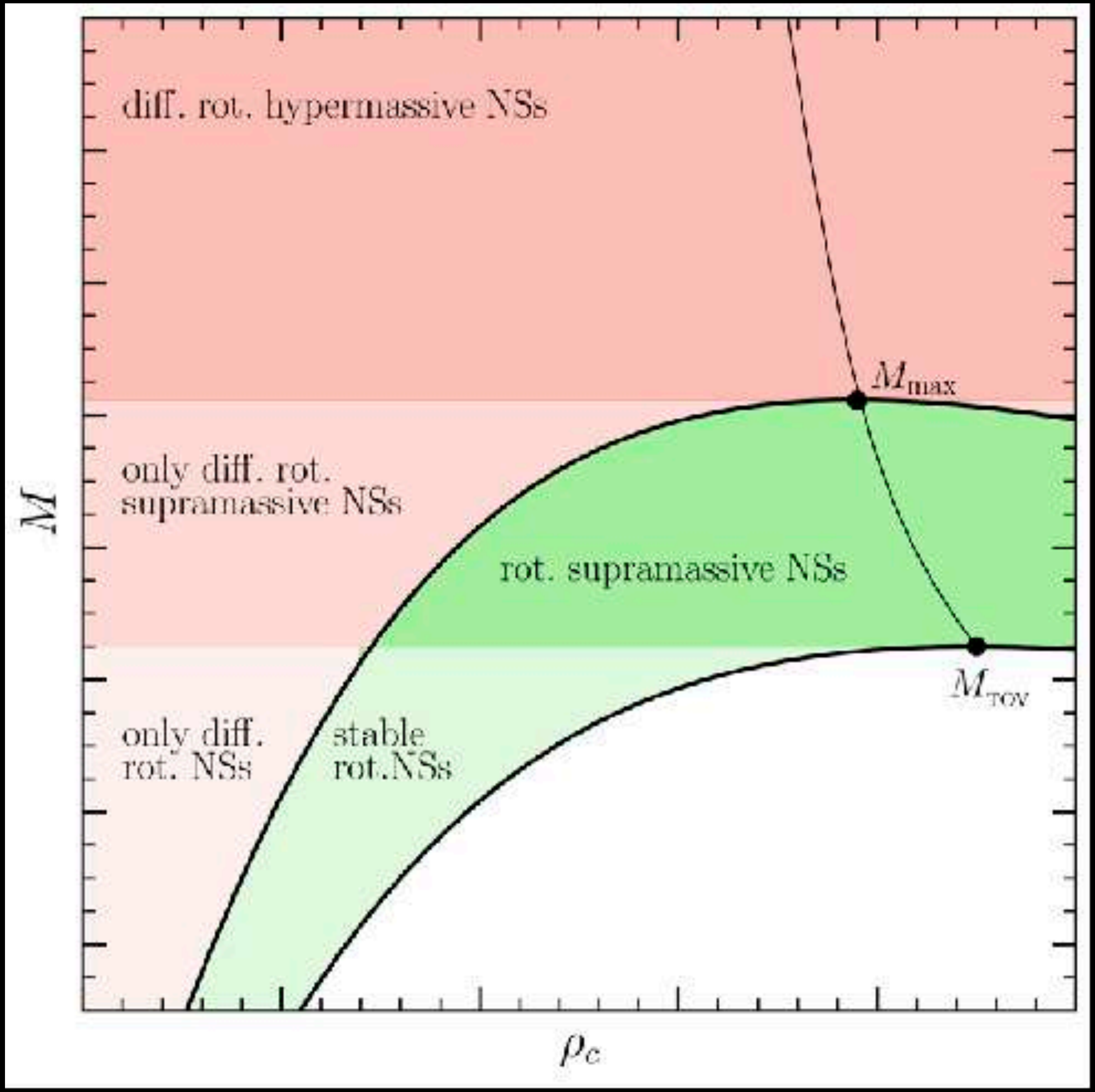


- **Green** region is for **uniformly** rotating equilibrium models.
- **Salmon** region is for **differentially** rotating equilibrium models.
- **Stability line** is simply extended in larger space (Weih+18)

Limits on the maximum mass

- The remnant of GW170817 was a hypermassive star, i.e. a differentially rotating object with initial **gravitational** mass:

$$M_1 + M_2 = 2.74^{+0.04}_{-0.01}\, M_\odot$$



- **Green** region is for **uniformly** rotating equilibrium models.
- **Salmon** region is for **differentially** rotating equilibrium models.
- **Supramassive** stars have:

$$M > M_{\rm TOV}$$

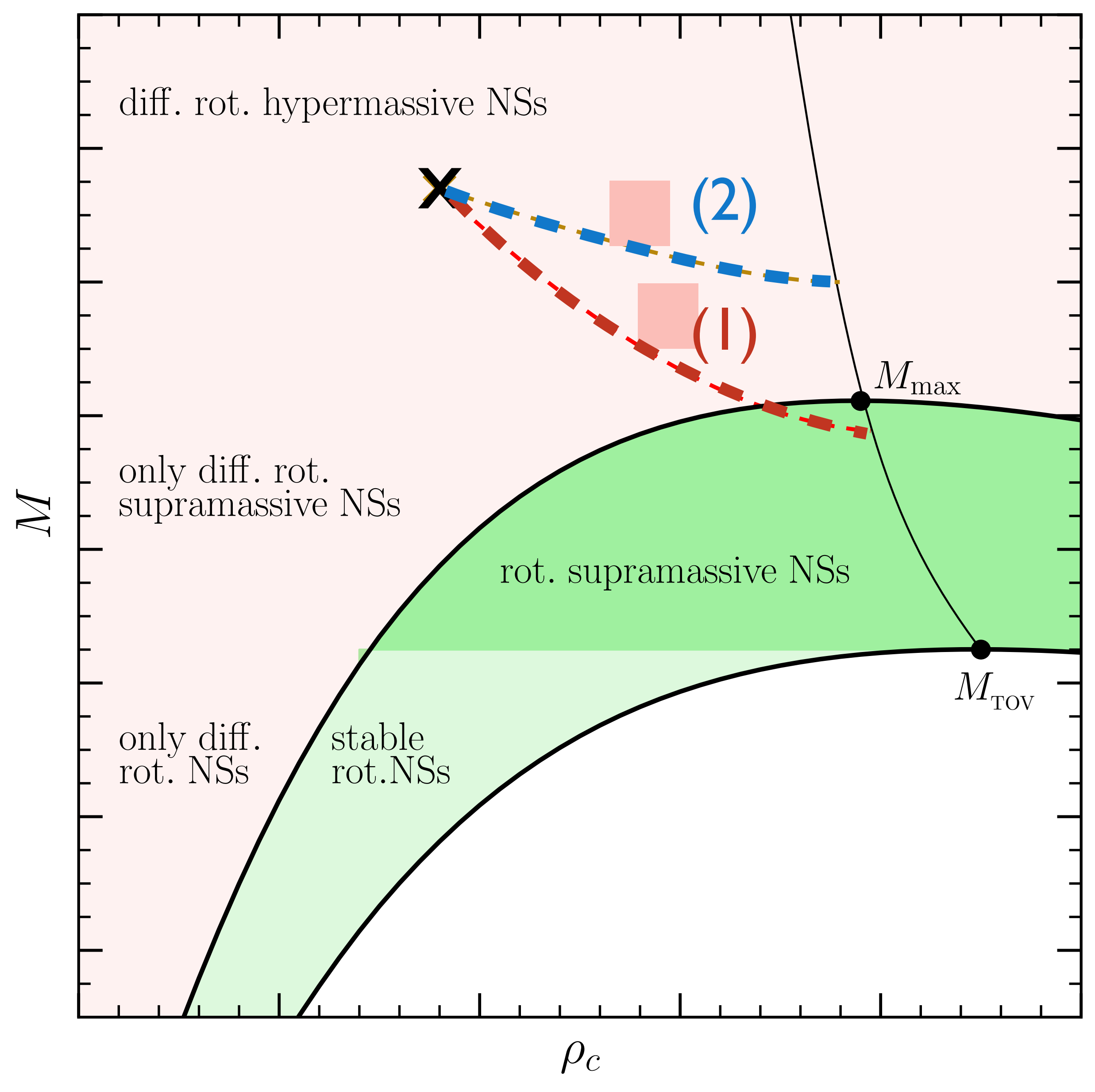
- **Hypermassive** stars have:

$$M > M_{\rm max}$$

Limits on the maximum mass

- GW170817 produced object **”X”**; GRB implies a BH has been formed: **”X”** followed two possible tracks: **fast (2)** and **slow (1)**

- It rapidly produced a BH when still **differentially** rotating (2)
- It lost differential rotation leading to a **uniformly** rotating core (1).
- (1) is much more likely because of large ejected mass (long lived).
- Final mass is near $M_{\rm max}$ and we know this is universal!



let’s recap…

- Consider **evolution track (1)**
- Use measured **gravitational mass** of GW170817
- Remove **rest-mass** deduced from kilonova emission (need conversion baryon/gravitational)
- Use **universal relations,** account for errors to obtain

pulsar timing

$$\boxed{2.01^{+0.04}_{-0.04}} \leq M_{\rm TOV}/M_{\odot} \leq \boxed{2.16^{+0.17}_{-0.15}}$$

GW170817; similar estimates by other groups (Margalit+ 2018, Shibata+ 2018, Ruiz+ 2018)

Tension on the maximum mass

Nathanail, Most, LR (2021)

- The recent detection of GW190814 has created a significant tension on the maximum mass

$$M_1 = 22.2 - 24.3\, M_\odot$$

$$M_2 = 2.50 - 2.67\, M_\odot$$

smallest BH or heaviest NS!

- If secondary in GW190814 was a NS, all previous results on the maximum mass are incorrect.
- No EM counterpart was observed with GW190814 and no estimates possible for ejected matter or timescale for survival.
- **How do we solve this tension?**

Tension on the maximum mass

- We can nevertheless explore impact of larger maximum mass, i.e., what changes in the previous picture if

$$M_{\rm TOV}/M_{\odot} \gtrsim 2.5\ ?$$

- In essence, this is a multi-dimensional parametric problem satisfying **conservation** of **rest-mass** and **gravitational mass.**
- Observations provide limits on **gravitational** and **ejected mass.**
- Numerical relativity simulations provide limits on **emitted GWs**
- All the rest is contained in **10 parameters** that need to be varied within suitable ranges.

Genetic algorithm

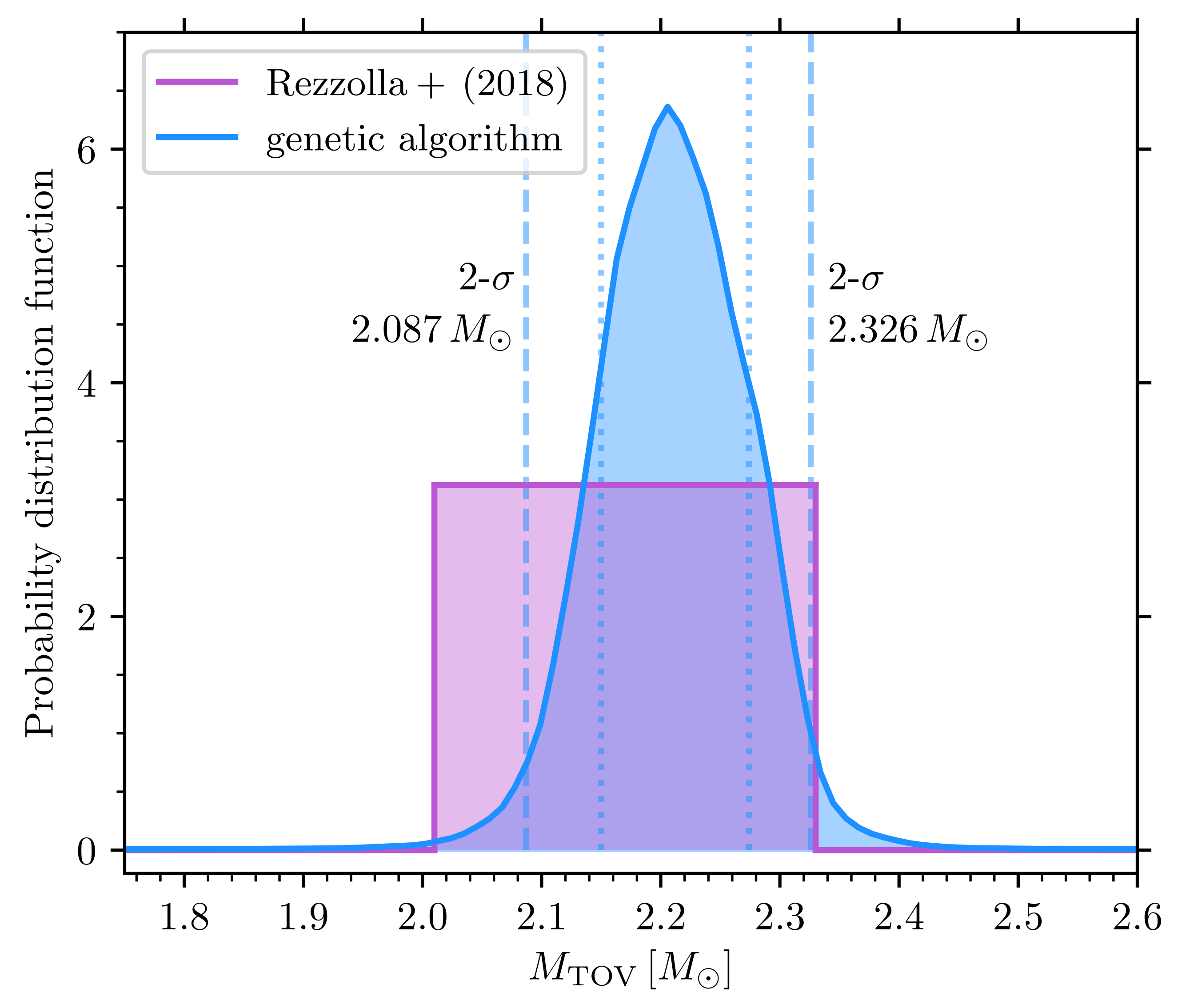
- A **genetic algorithm** is used to sample through the parameter space of the 10 free parameters.

- The algorithm reflects genetic adaptation: given a mutation (i.e. change of parameters) it will be adopted if it provides a better fit to data.

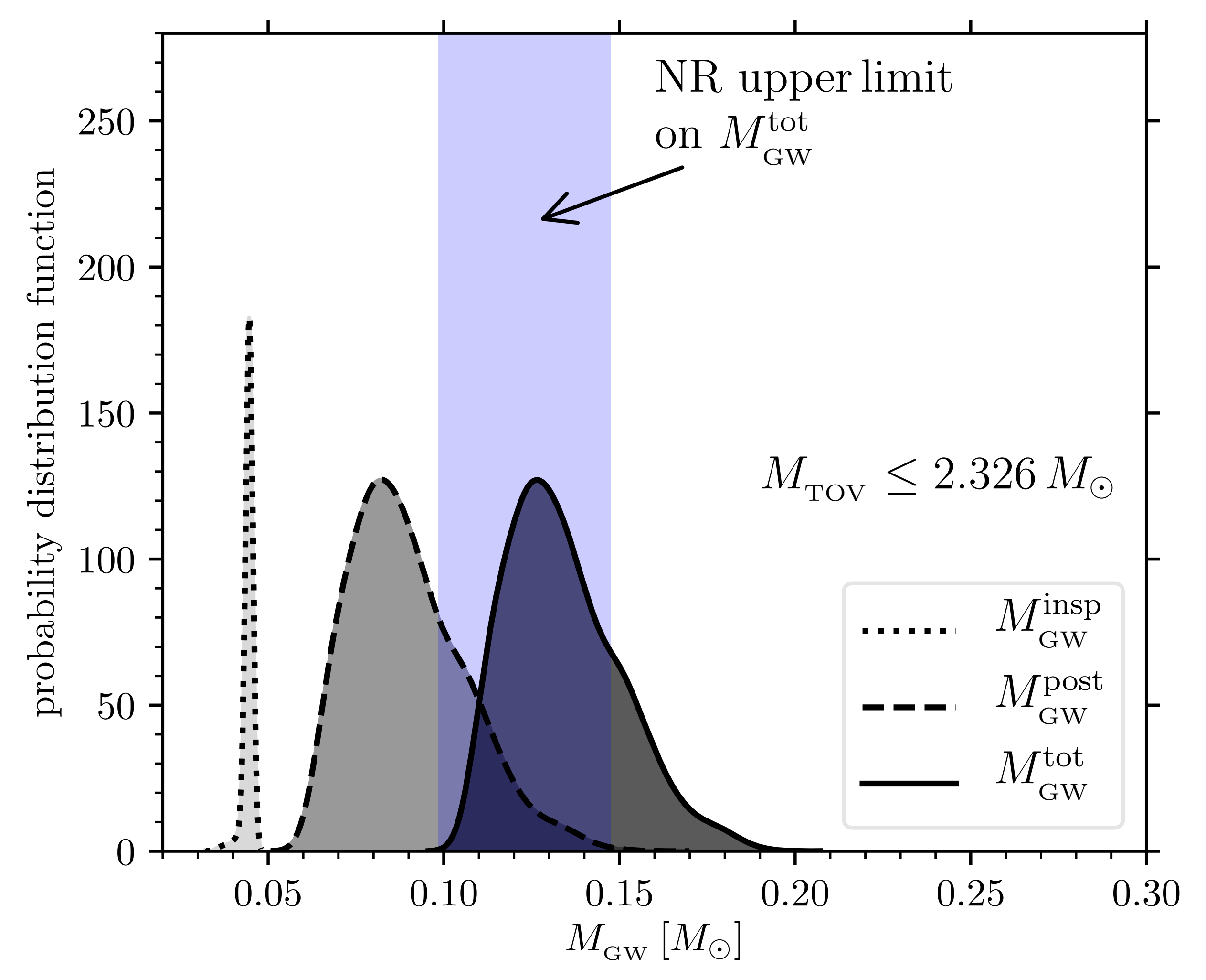
- Consider first previous estimate:

$$M_{\rm TOV}/M_\odot \lesssim 2.3$$

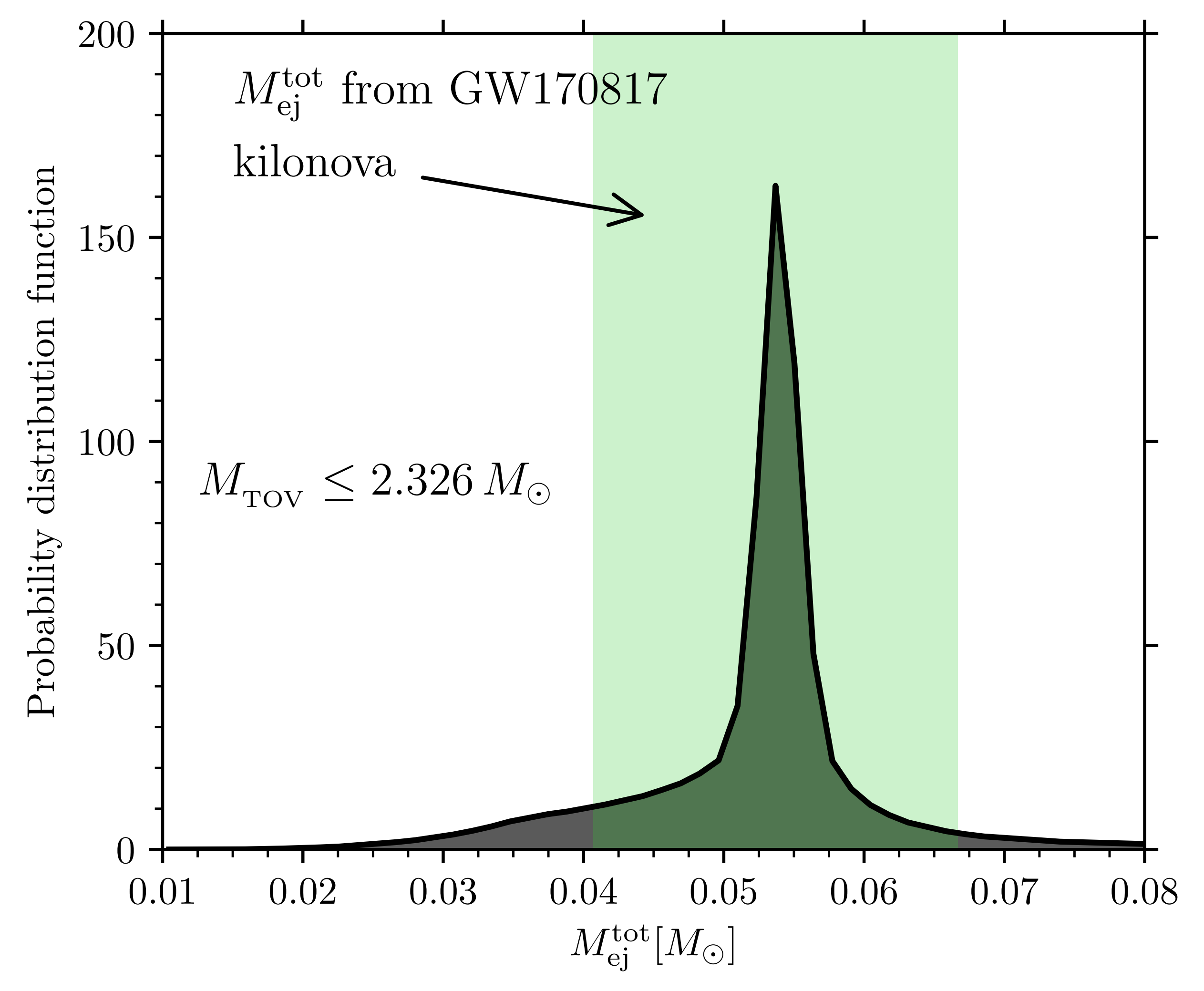
$$M_{\rm TOV}/M_\odot \leq 2.16^{+0.17}_{-0.15}$$



First hypothesis: $M_{\text{TOV}}/M_{\odot} \lesssim 2.3$

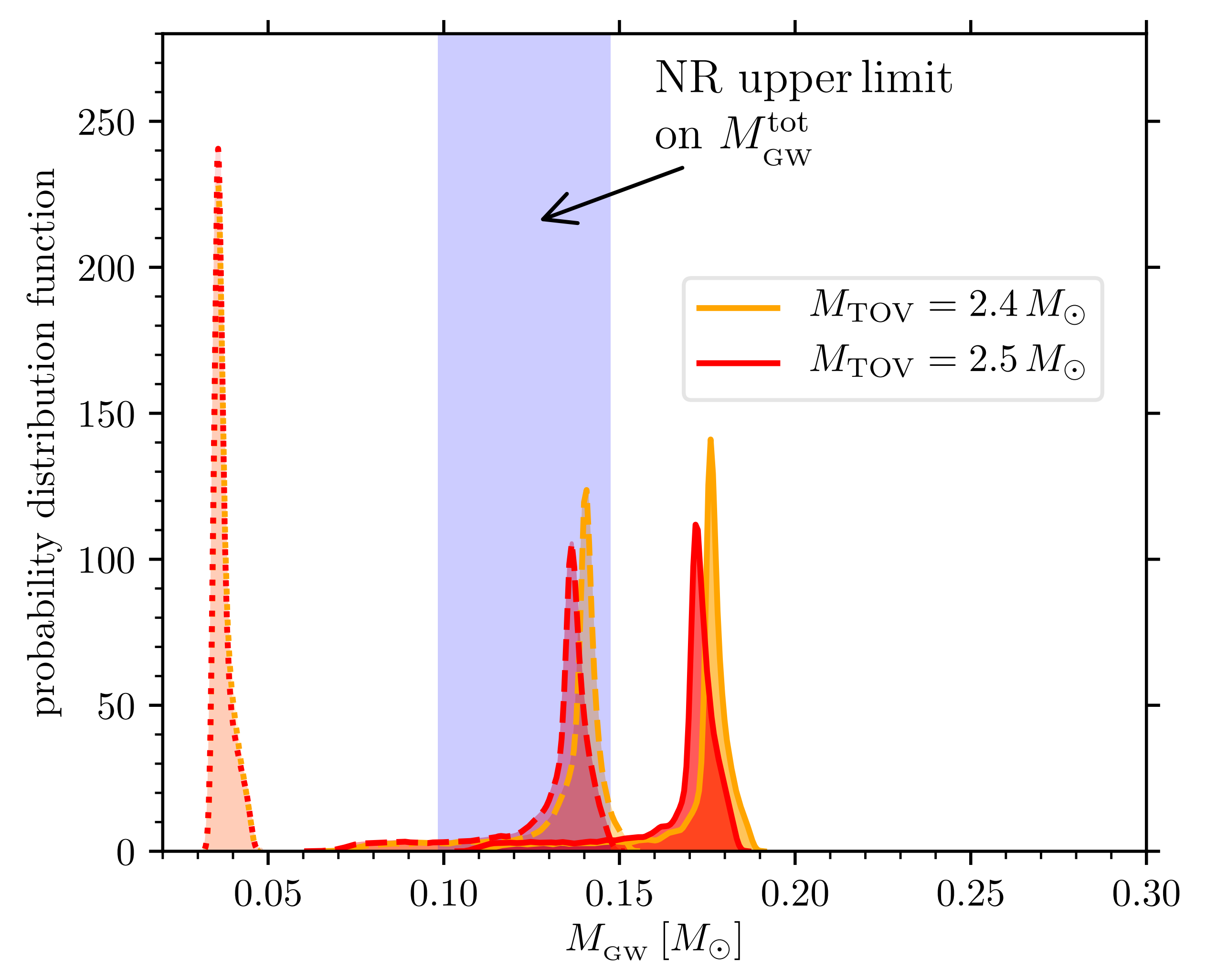


- Total mass emitted in GWs is in perfect **agreement** with predictions from numerical relativity

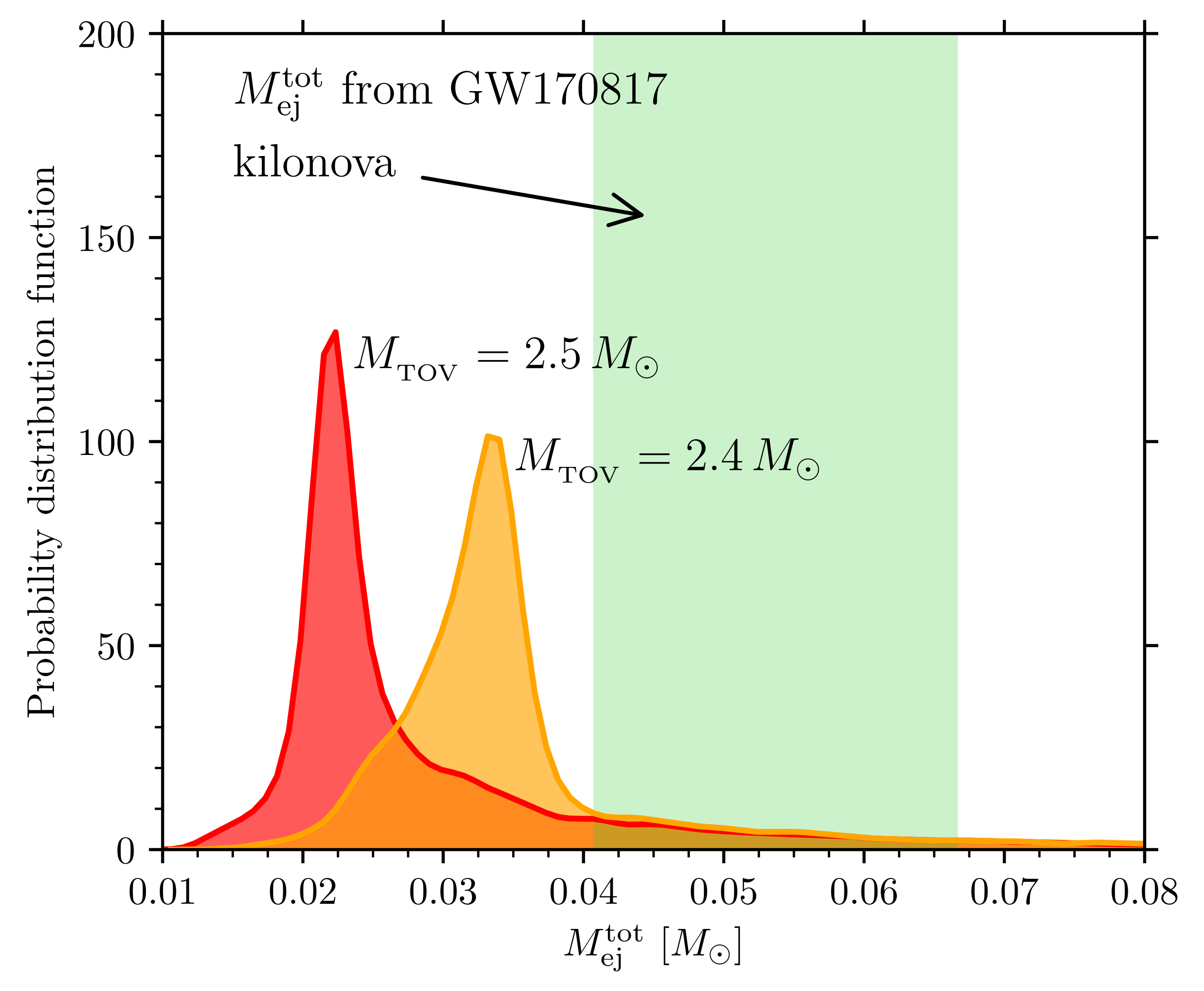


- Total mass ejected is in perfect **agreement** with predictions from kilonova signal

Second hypothesis: $M_{\rm TOV}/M_\odot \gtrsim 2.5$



- Total mass emitted in GWs is **much larger** than predicted from simulations;
- Mismatch becomes worse with larger masses



- Total mass ejected is in perfect **much smaller** than observed from kilonova signal.

Tension on the maximum mass

Nathanail, Most, LR (2020)

- The recent detection of GW190814 has created a significant tension on the maximum mass

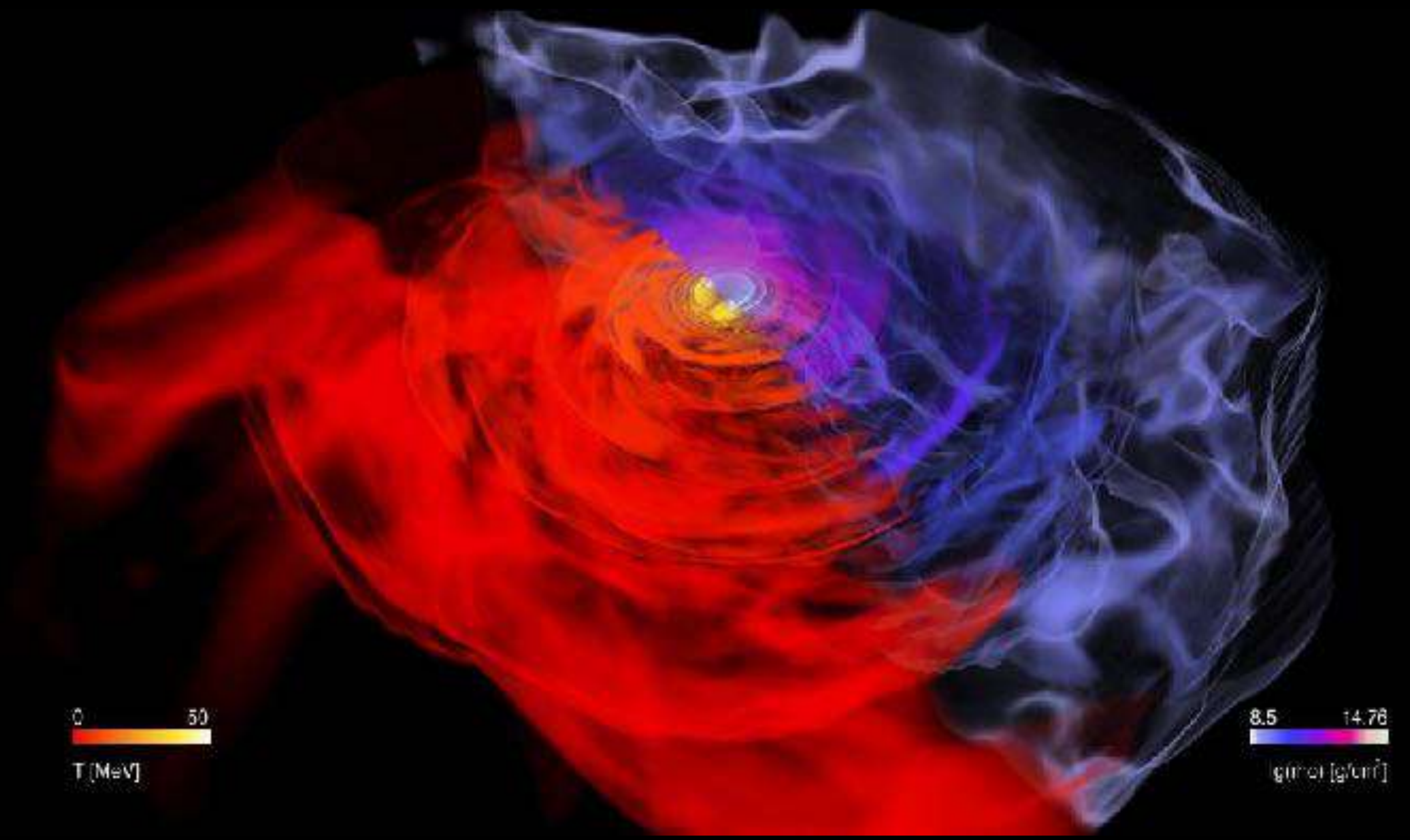
$$M_1 = 22.2 - 24.3\, M_\odot$$

$$M_2 = 2.50 - 2.67\, M_\odot$$

smallest BH or heaviest NS!

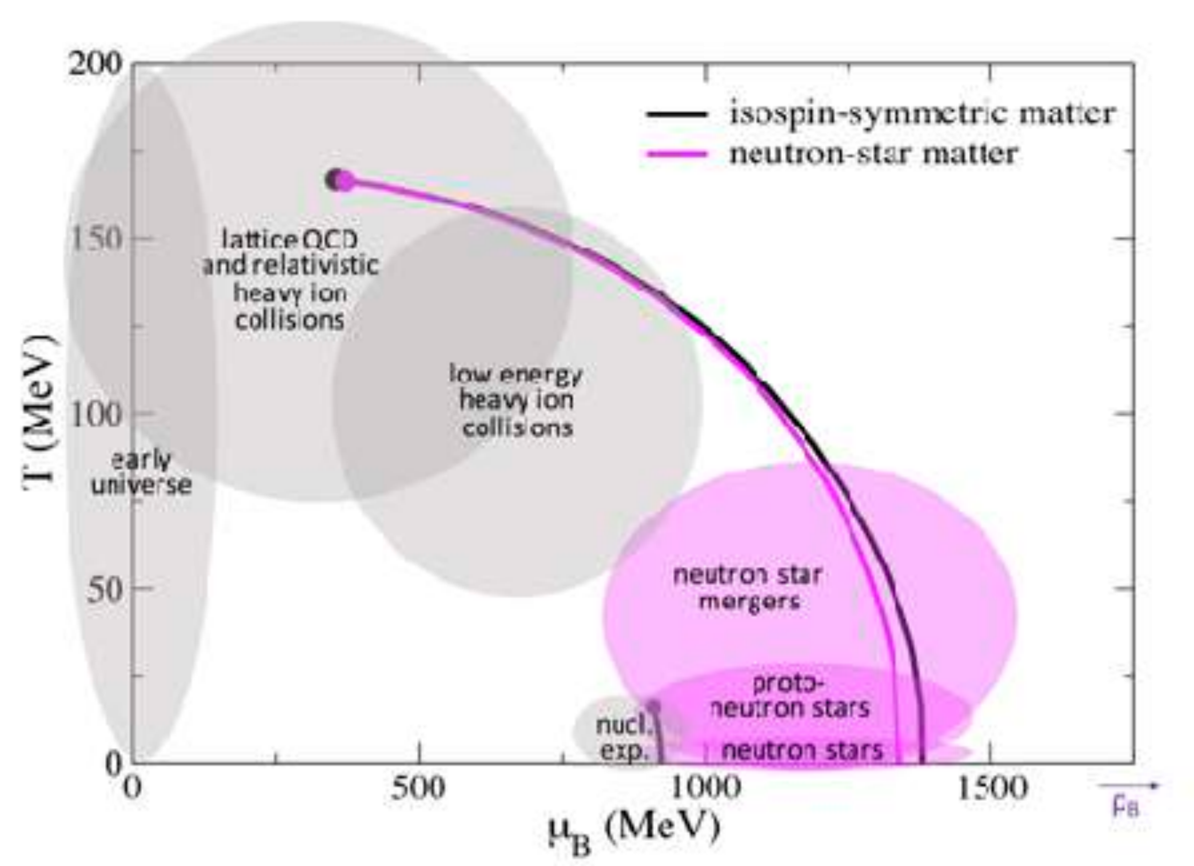
- If secondary in GW190814 was a NS, all previous considerations are incorrect.
- No EM counterpart was observed with GW190814 and no estimates possible on ejected matter or timescale for survival.
- **How do we solve this tension?**
- Solution: secondary in GW190814 **was a BH at merger** but could have been a NS before

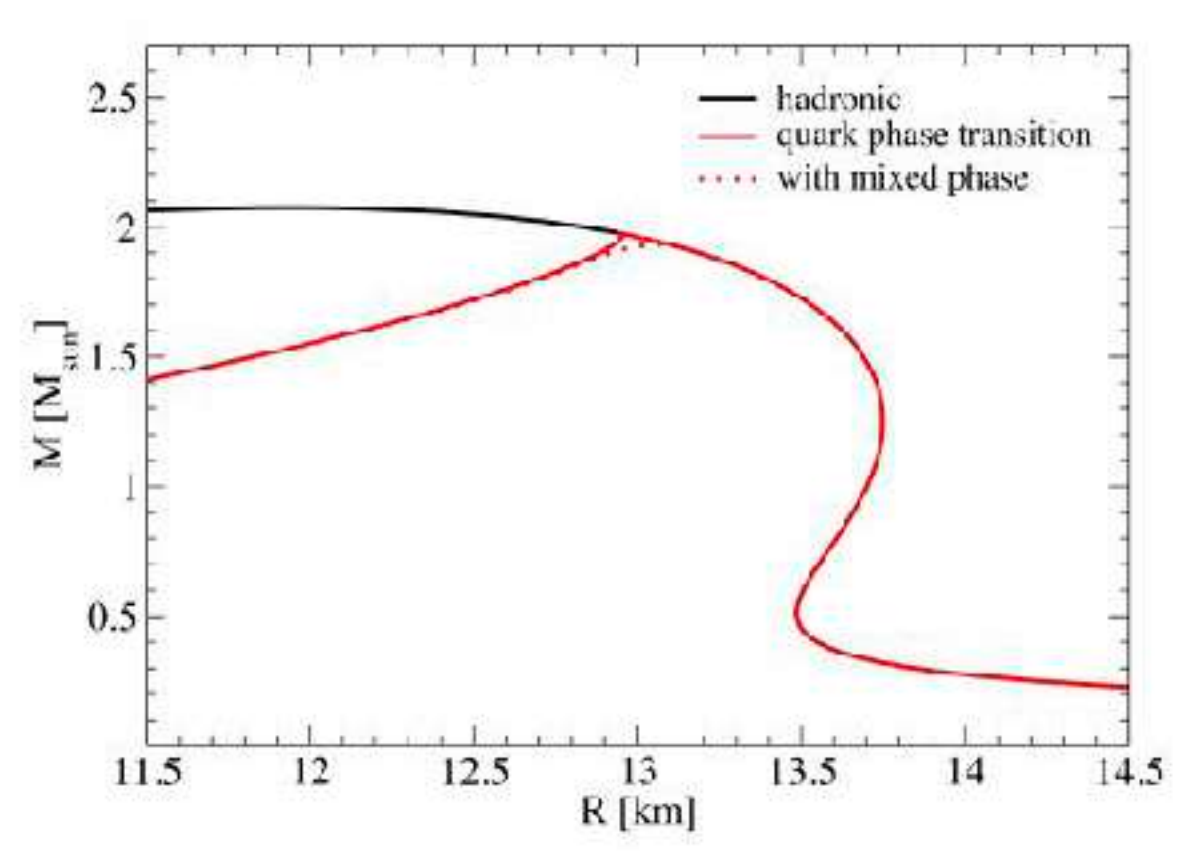
Phase transitions and their signatures



Most, Papenfort, Dexheimer, Hanauske, Schramm, Stoecker, LR (2019)
Weih, Hanauske, LR (2020)
Tootle, Ecker, Topolski, Demircik, Järvinen, LR (2022)

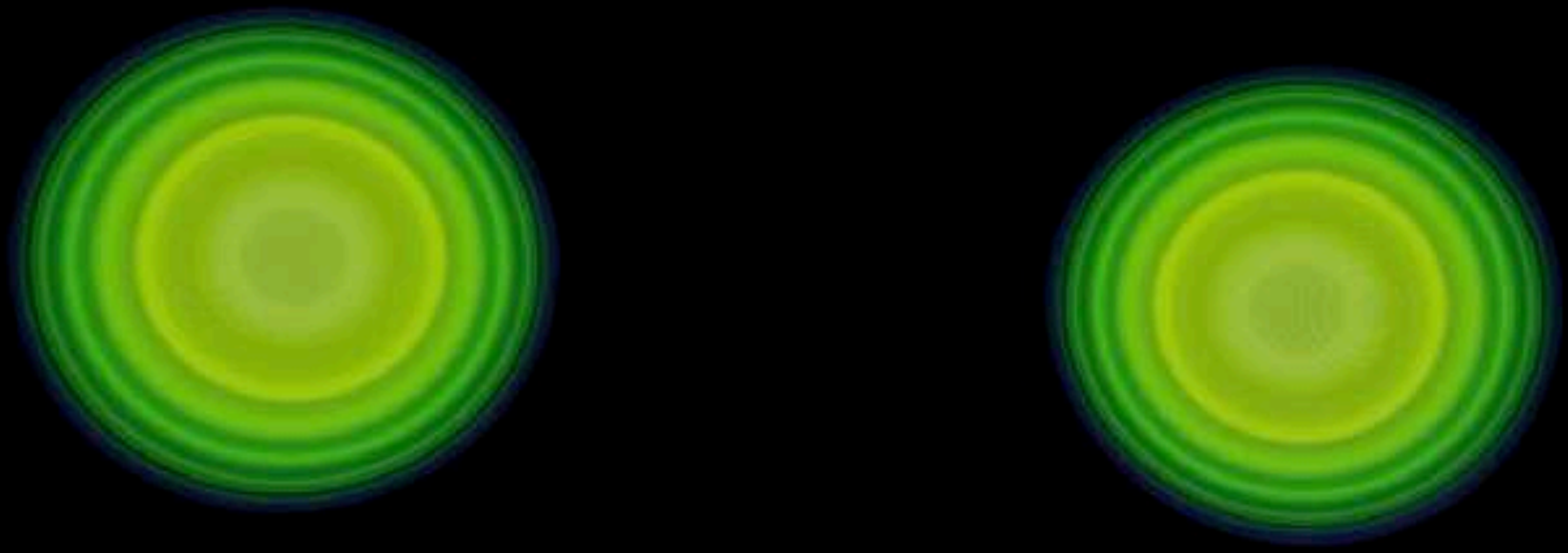
- **Isolated** neutron stars probe a small fraction of phase diagram.
- Neutron-star **binary** mergers reach temperatures up to 80 MeV and probe regions complementary to experiments.

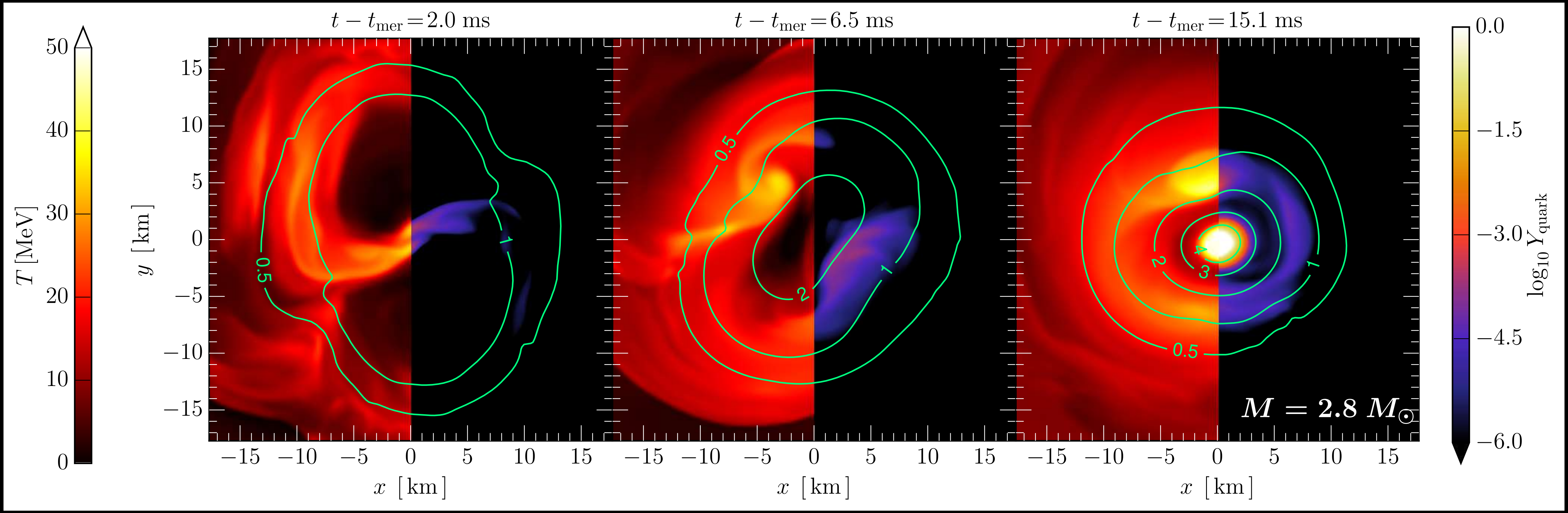




- Considered EOS based on Chiral Mean Field (CMF) model, based on a nonlinear SU(3) sigma model.
- Appearance of quarks can be introduced naturally.

Animations: Weih, Most, LR

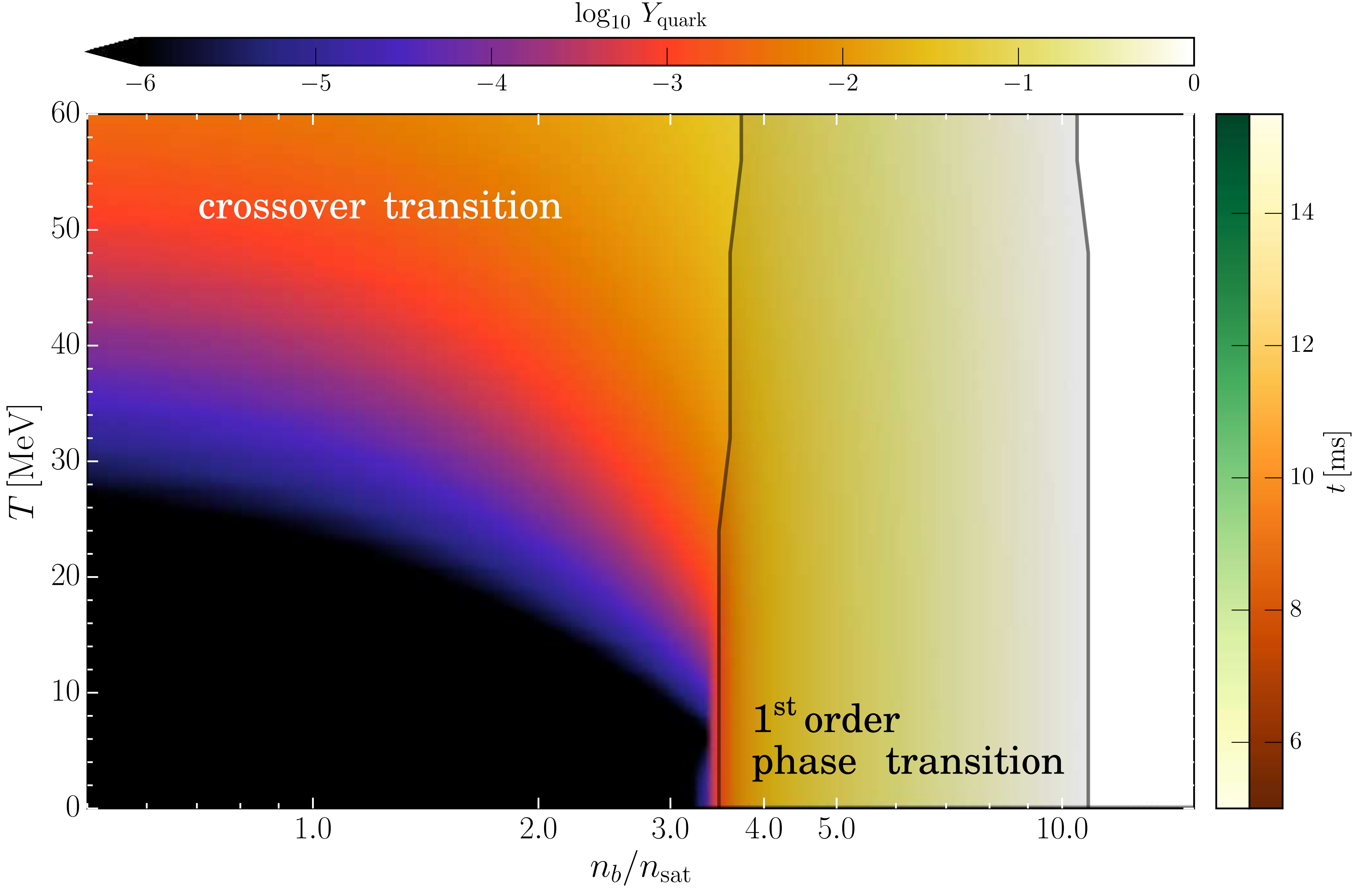




Quarks appear at sufficiently large temperatures and densities.

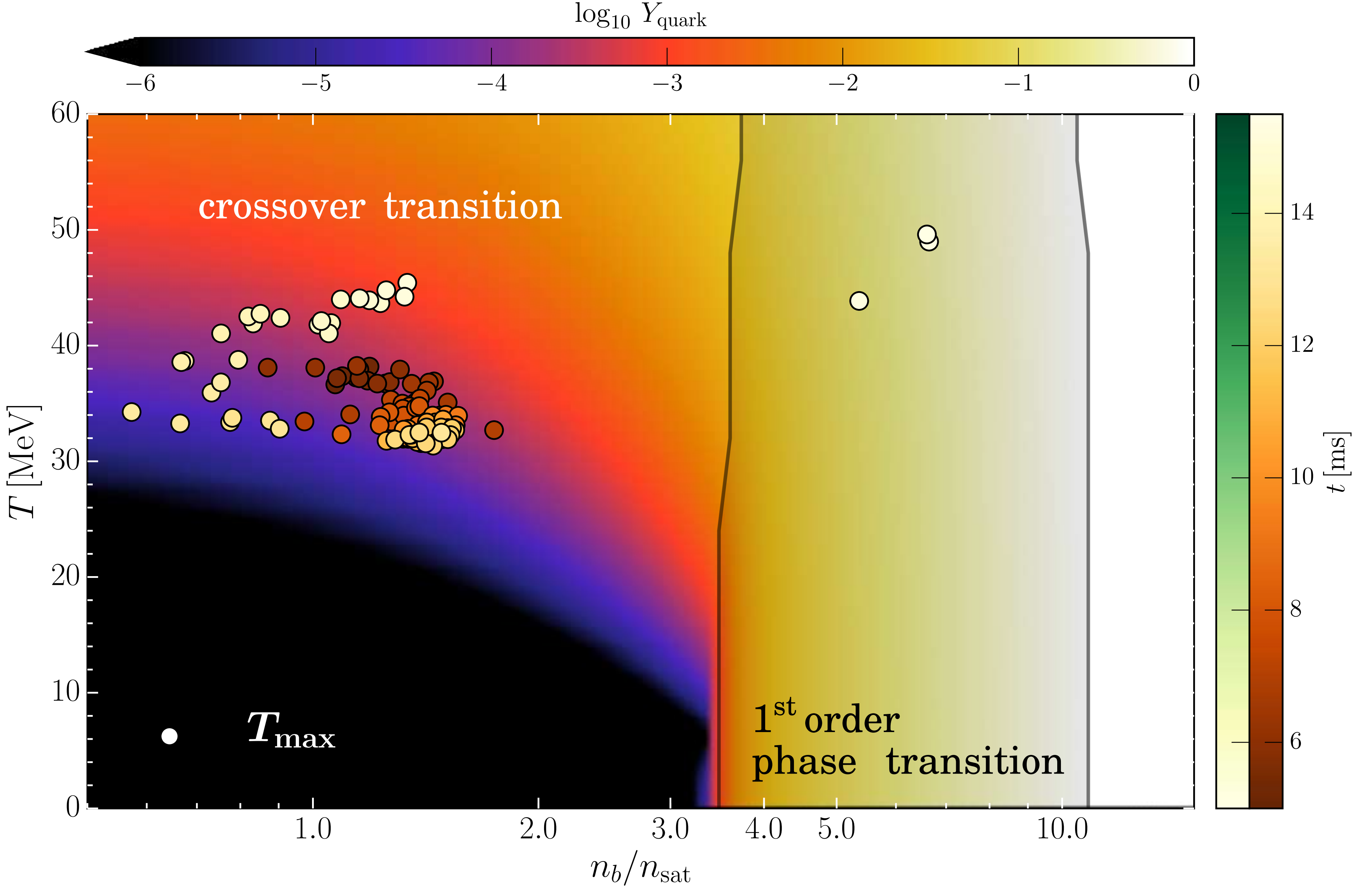
When this happens the EOS is considerably softened and a BH produced.

Comparing with the phase diagram



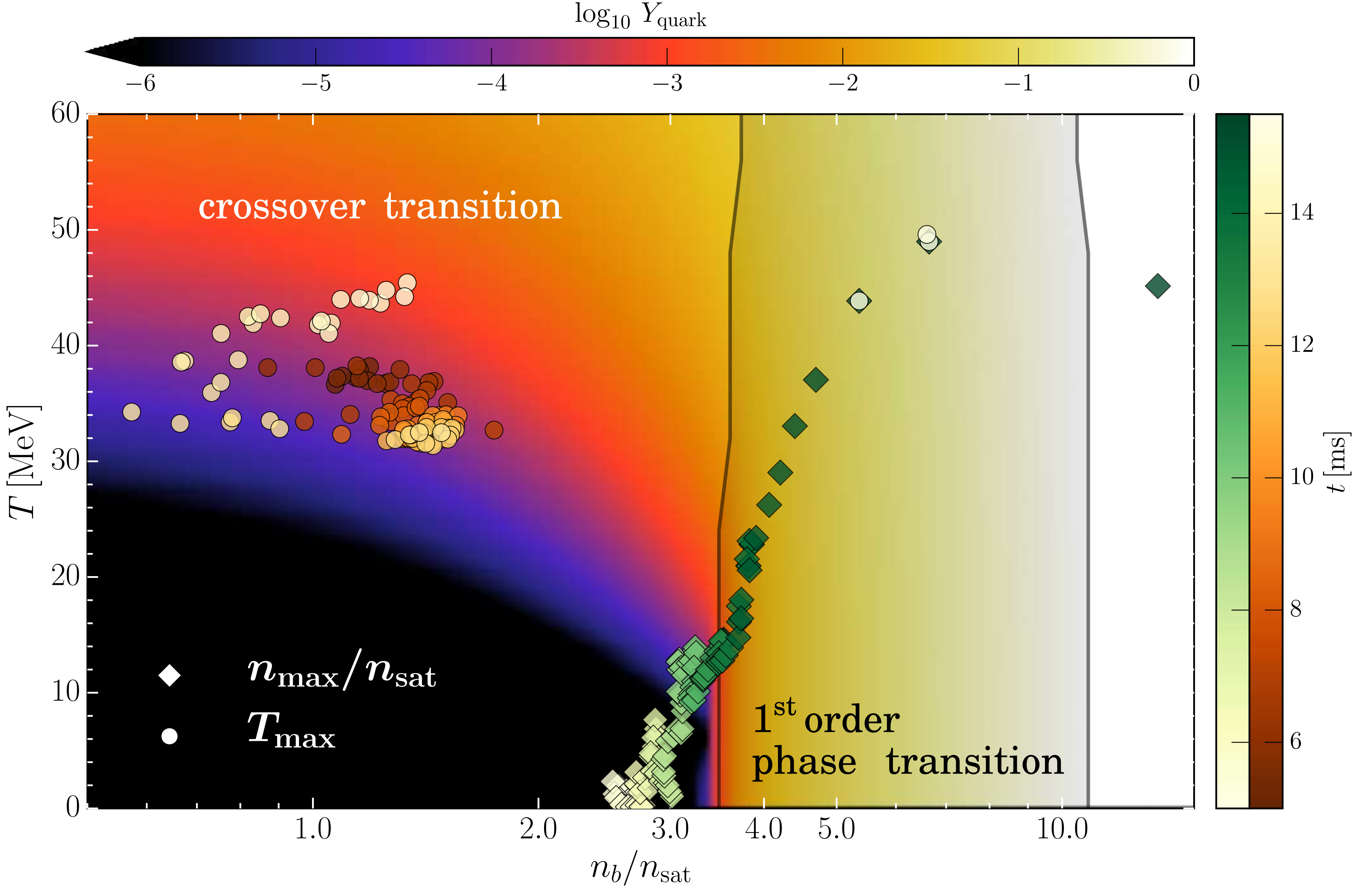
- Phase diagram with quark fraction

Comparing with the phase diagram



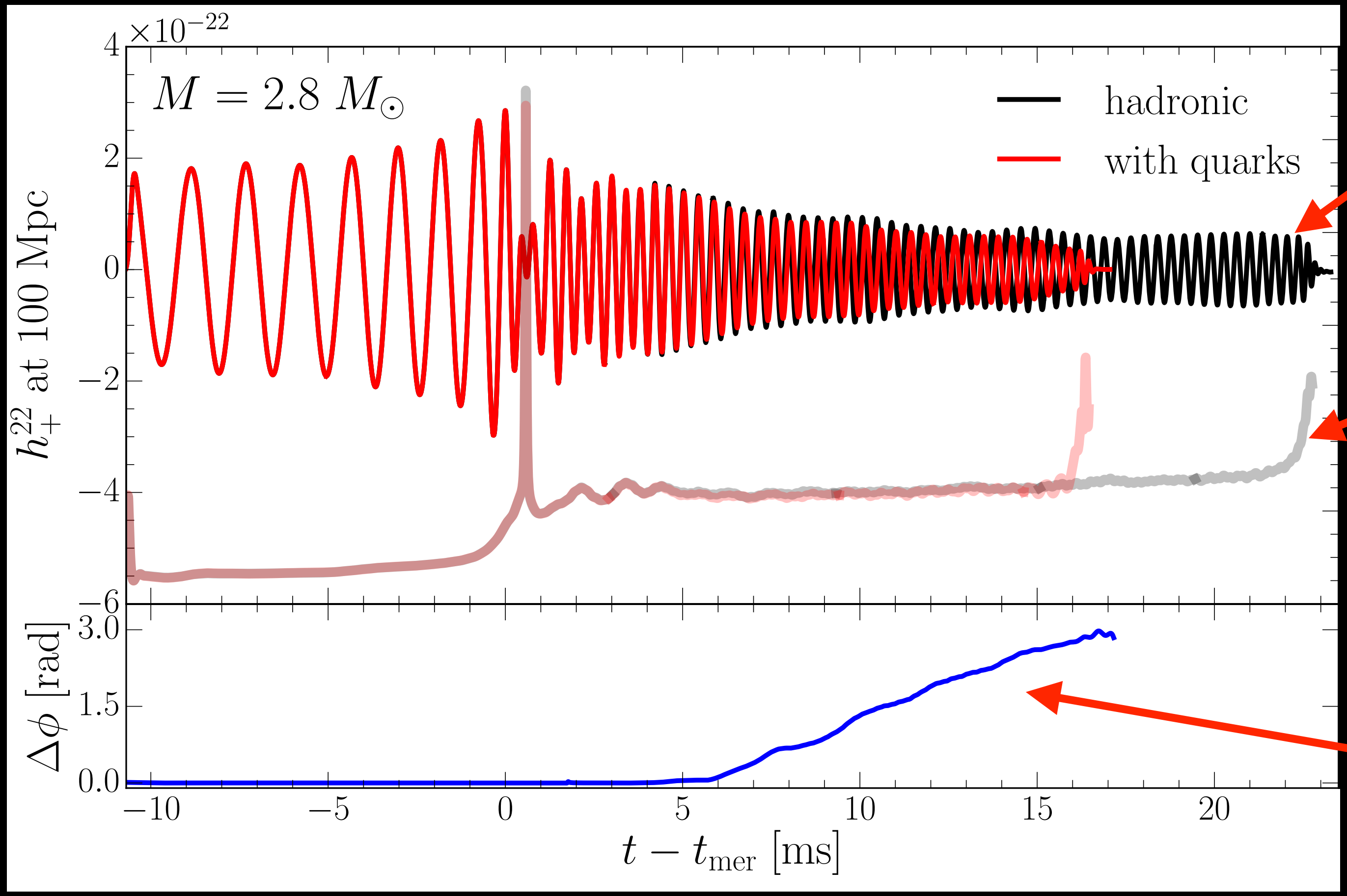
- Phase diagram with quark fraction
- Circles show the position in the diagram of the maximum temperature as a function of time

Comparing with the phase diagram



- Reported are the evolution of the max. temperature and density.
- Quarks appear already early on, but only in small fractions.
- Once sufficient density is reached, a full phase transition takes place.

Gravitational-wave emission



waveforms

GW frequencies

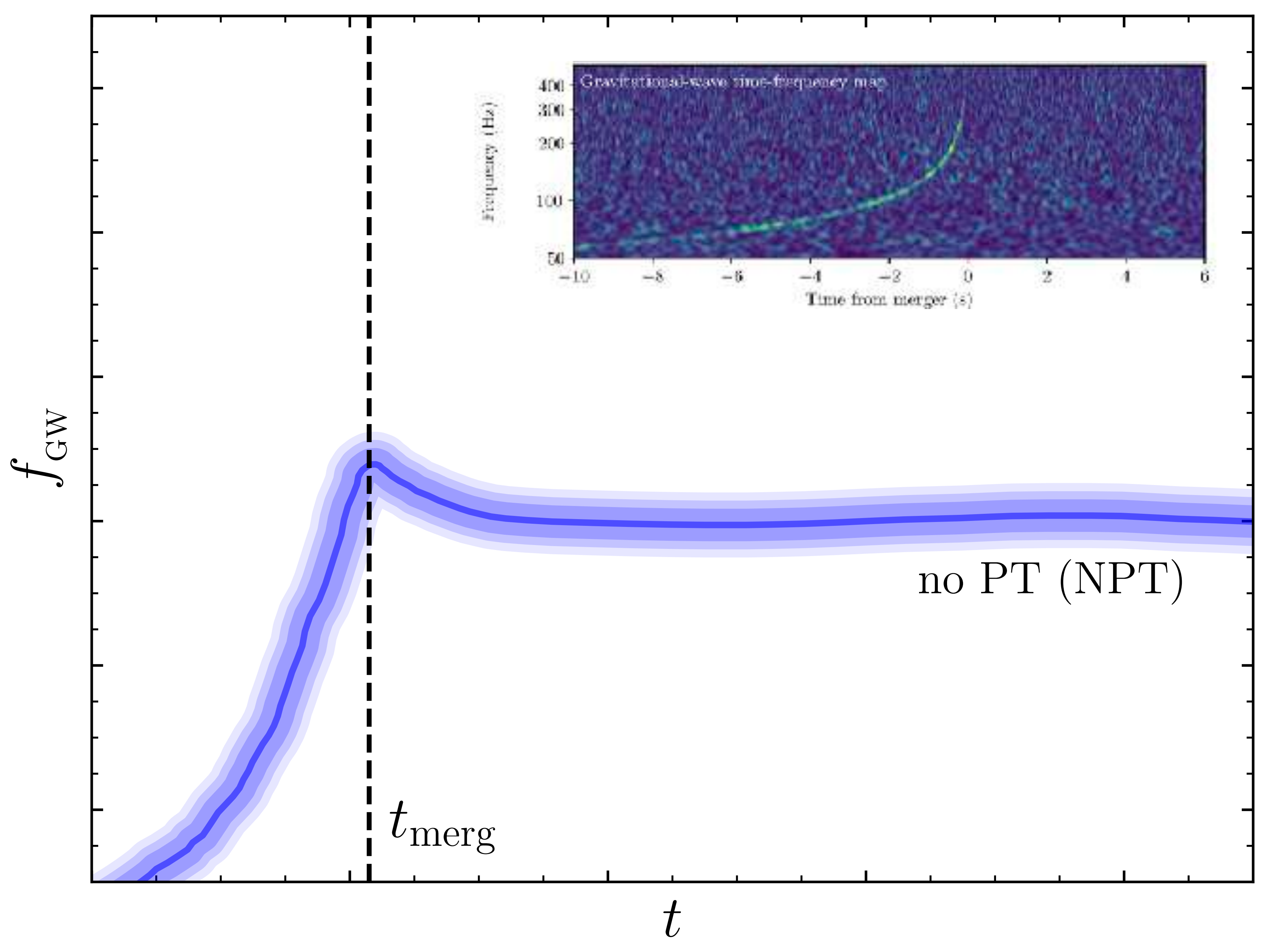
phase difference

- After ~ 5 ms, quark fraction is large enough to change quadrupole moment and yield differences in the waveforms.
- Sudden softening of the phase transition leads to collapse and **large difference** in phase evolution.

Observing mismatch between **inspiral** (fully hadronic) and **post-merger** (phase transition): clear **signature** of a **PT**

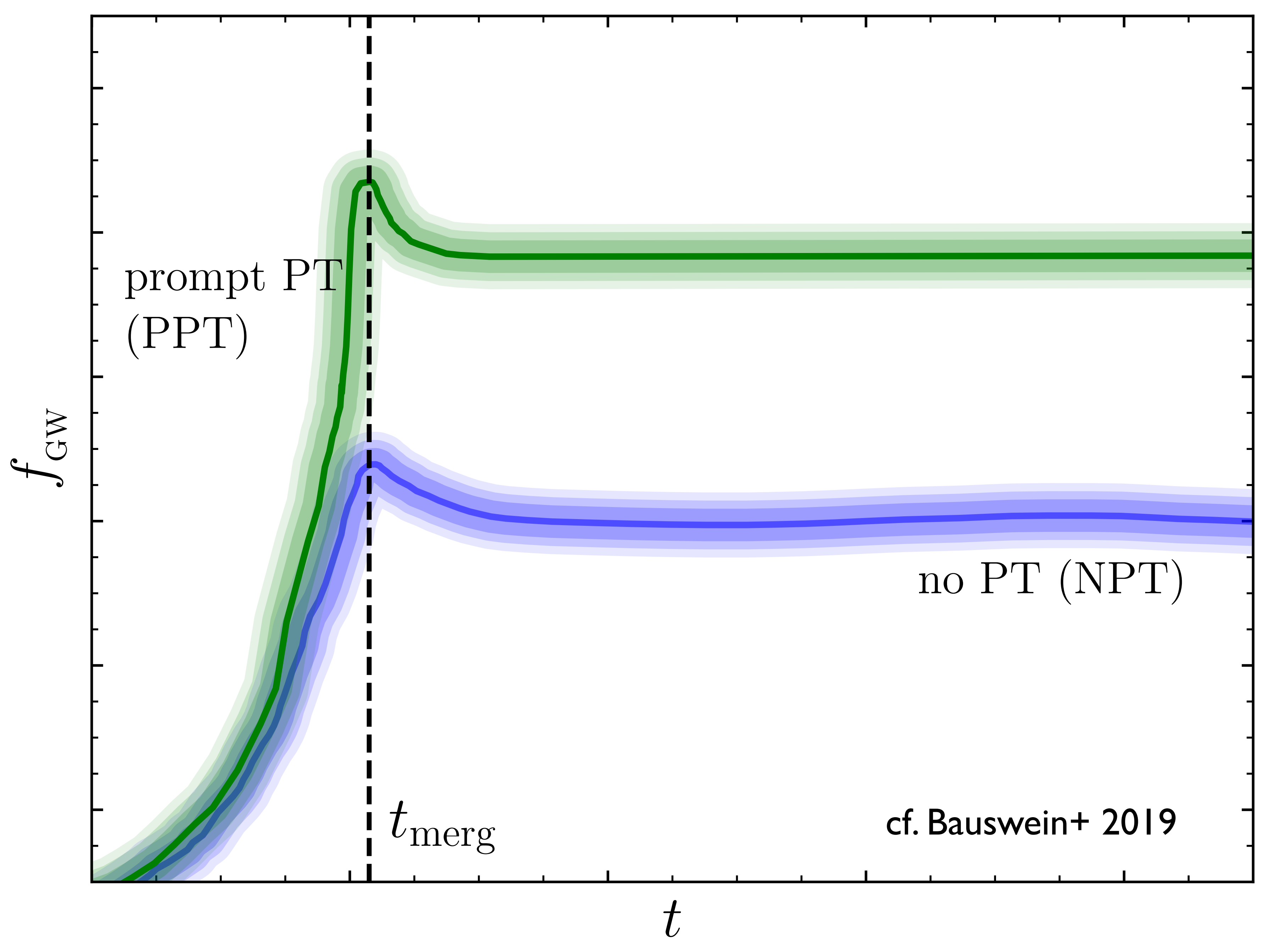
A more comprehensive picture

We have recently added another possible scenario for a post-merger **PT,** which completes the picture of possible scenarios (Weih, Hanauske, LR 2020).



A more comprehensive picture

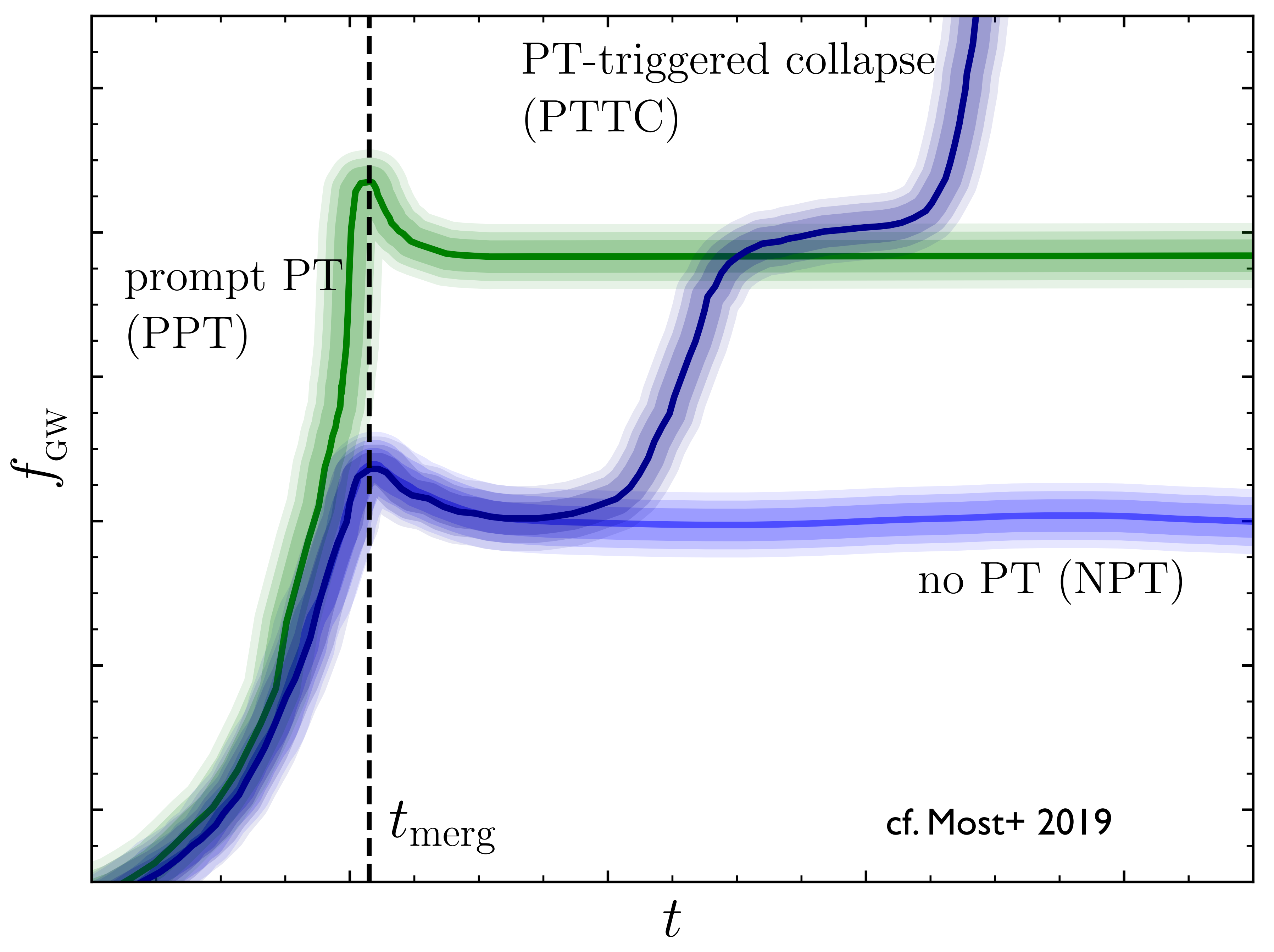
We have recently added another possible scenario for a post-merger **PT,** which completes the picture of possible scenarios (Weih+, 1912.09340).



cf. Bauswein+ 2019

A more comprehensive picture

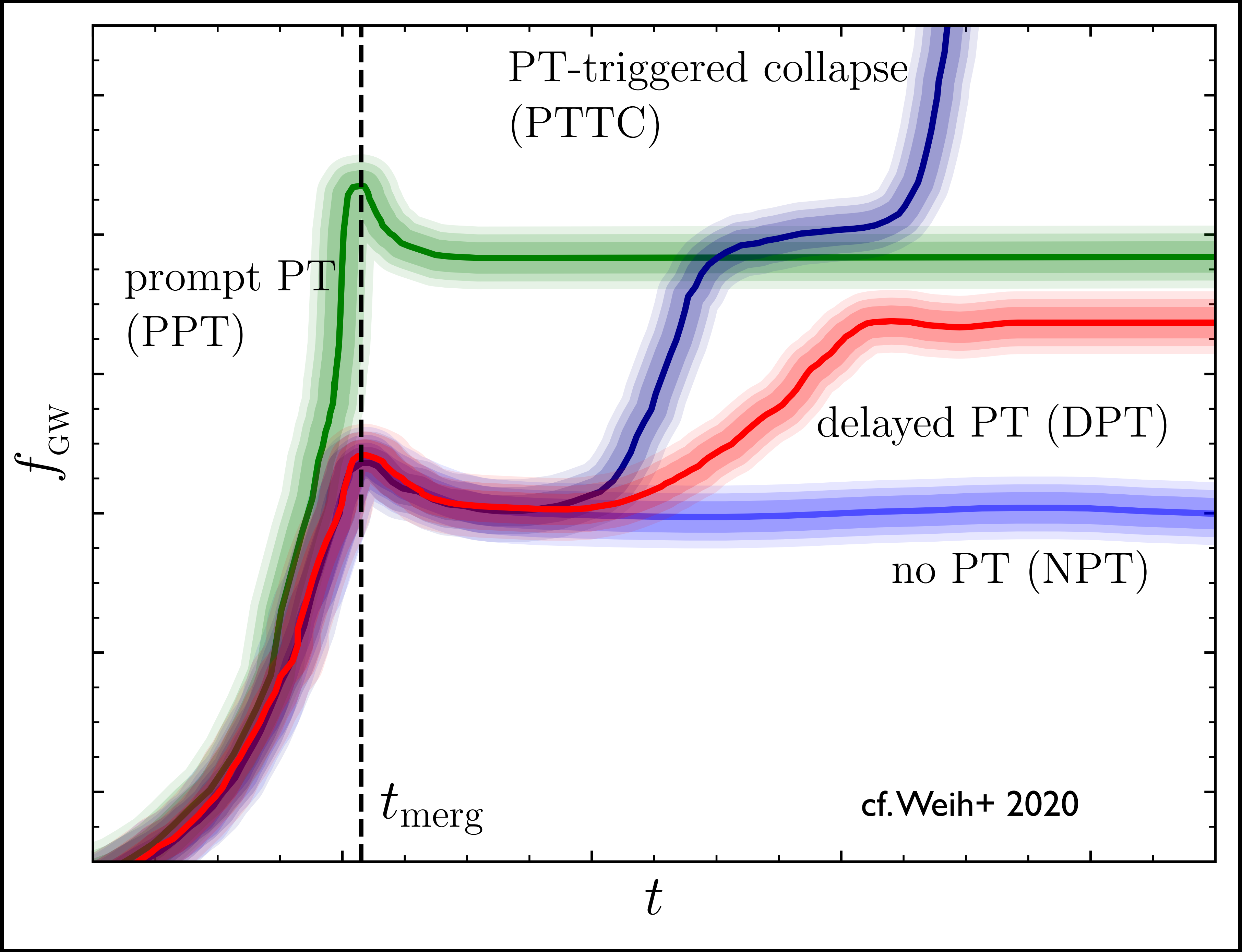
We have recently added another possible scenario for a post-merger **PT,** which completes the picture of possible scenarios (Weih+, 1912.09340).



cf. Most+ 2019

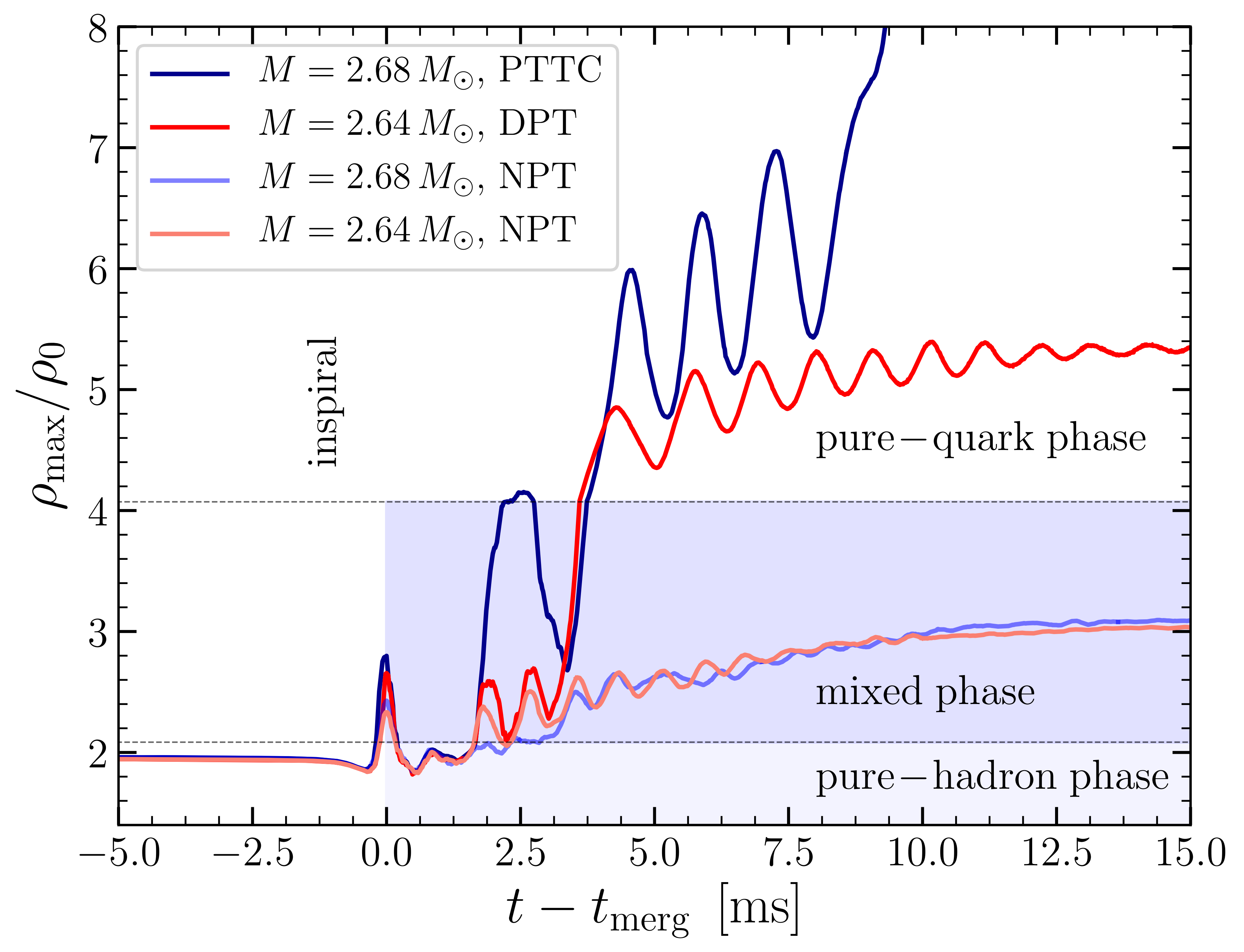
A more comprehensive picture

We have recently added another possible scenario for a post-merger **PT**, which completes the picture of possible scenarios (Weih+, 1912.09340).



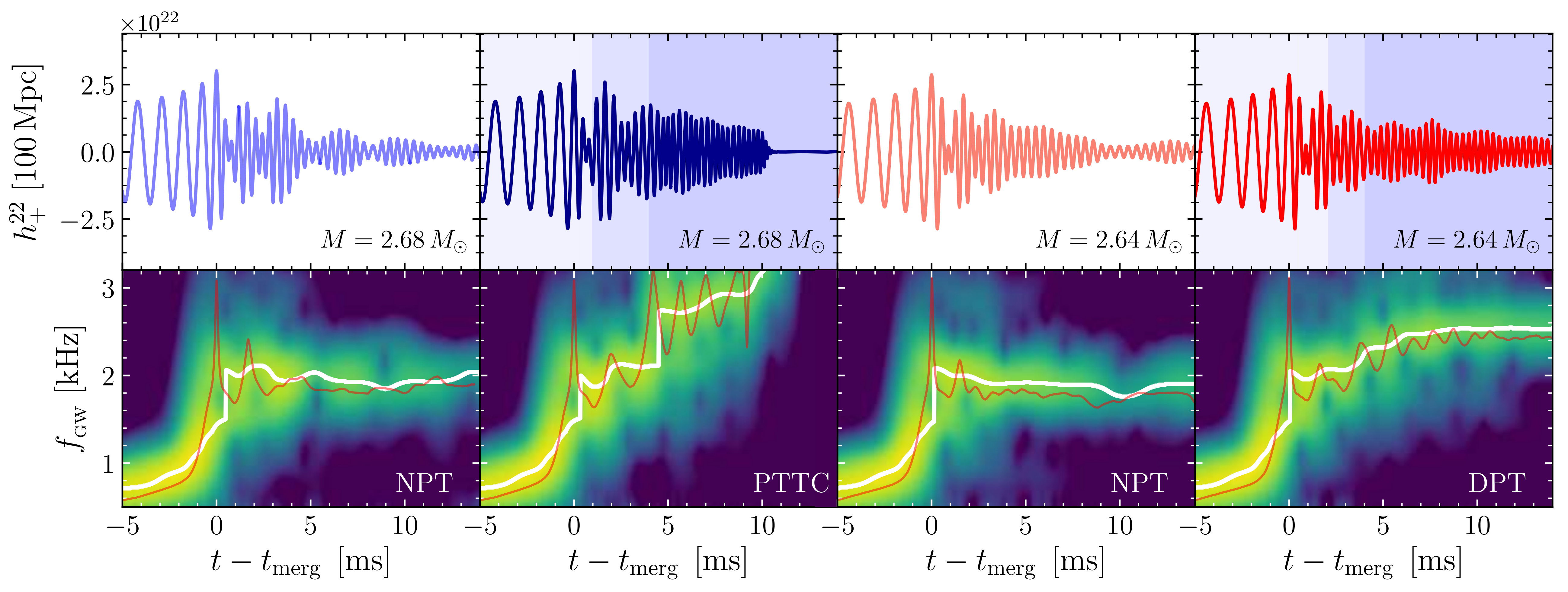
A more comprehensive picture

Best understood in terms of the evolution of the normalise maximum rest-mass density: $\rho_{\max}/\rho_0$



A more comprehensive picture

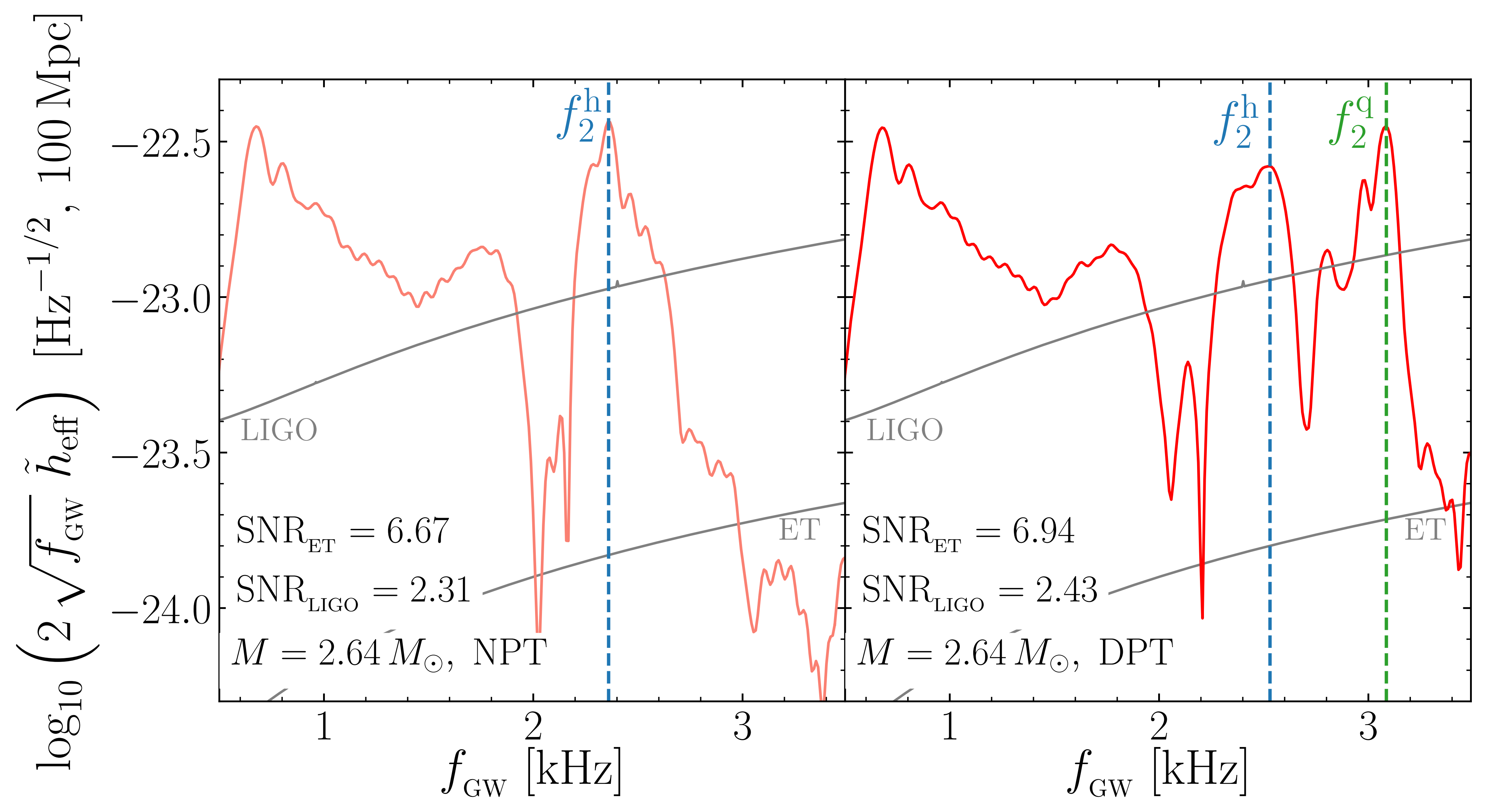
Different signatures are also quite transparent when shown in terms of the gravitational waves and their spectrograms.



Importance of **DPT** is that it leads to **two** different "stable" f_2 **frequencies** that are easily distinguishable in the PSD

A more comprehensive picture

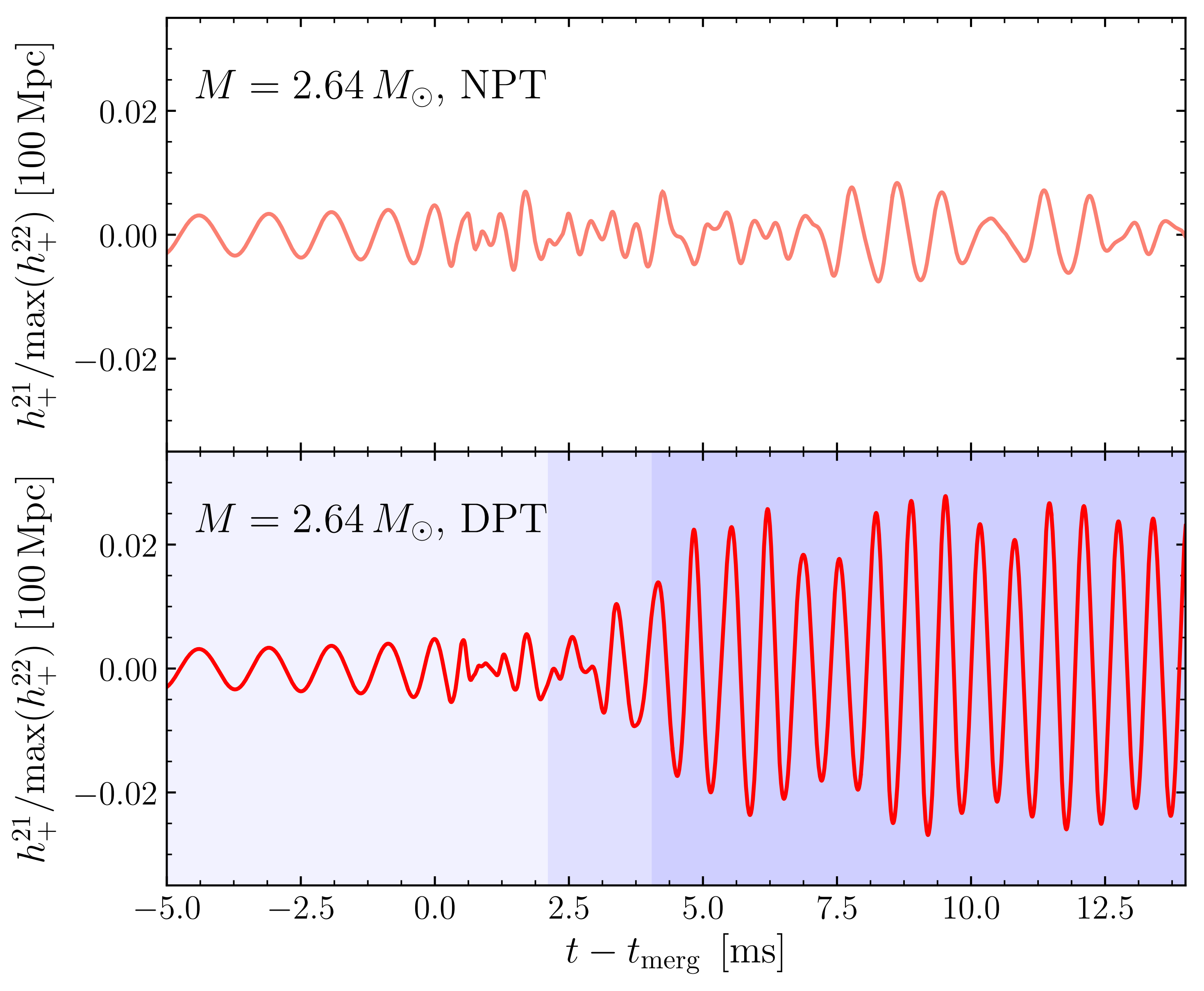
Different signatures are also quite transparent when shown in terms of the gravitational waves and their spectrograms.



Importance of **DPT** is that it leads to **two** different "stable" f_2 **frequencies** that are easily distinguishable in the PSD

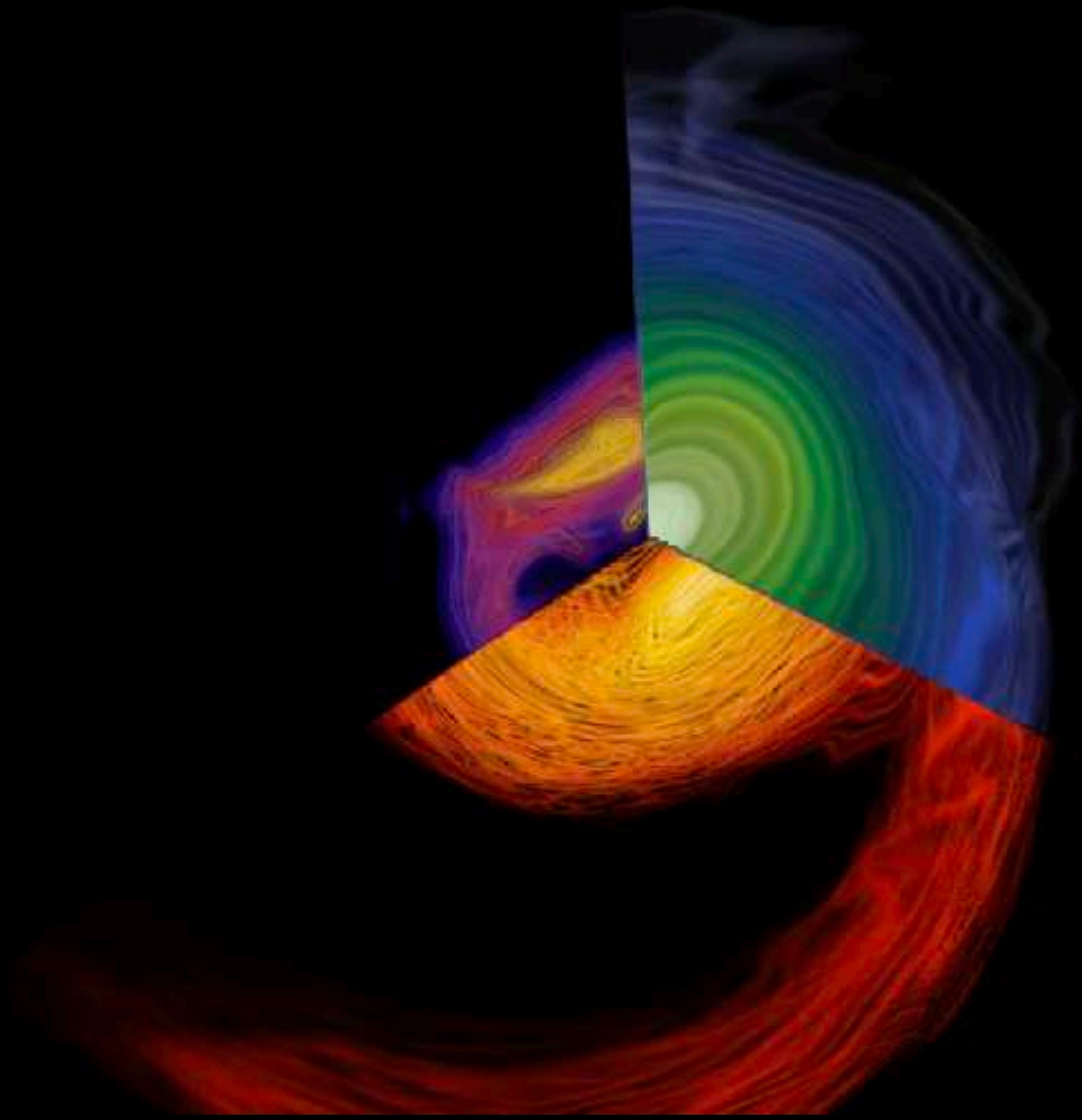
A more comprehensive picture

Another signatures is appearance of an $\ell = 2, m = 1$ mode



The mode is triggered by the PT and the non-axisymmetric deformations it produces.

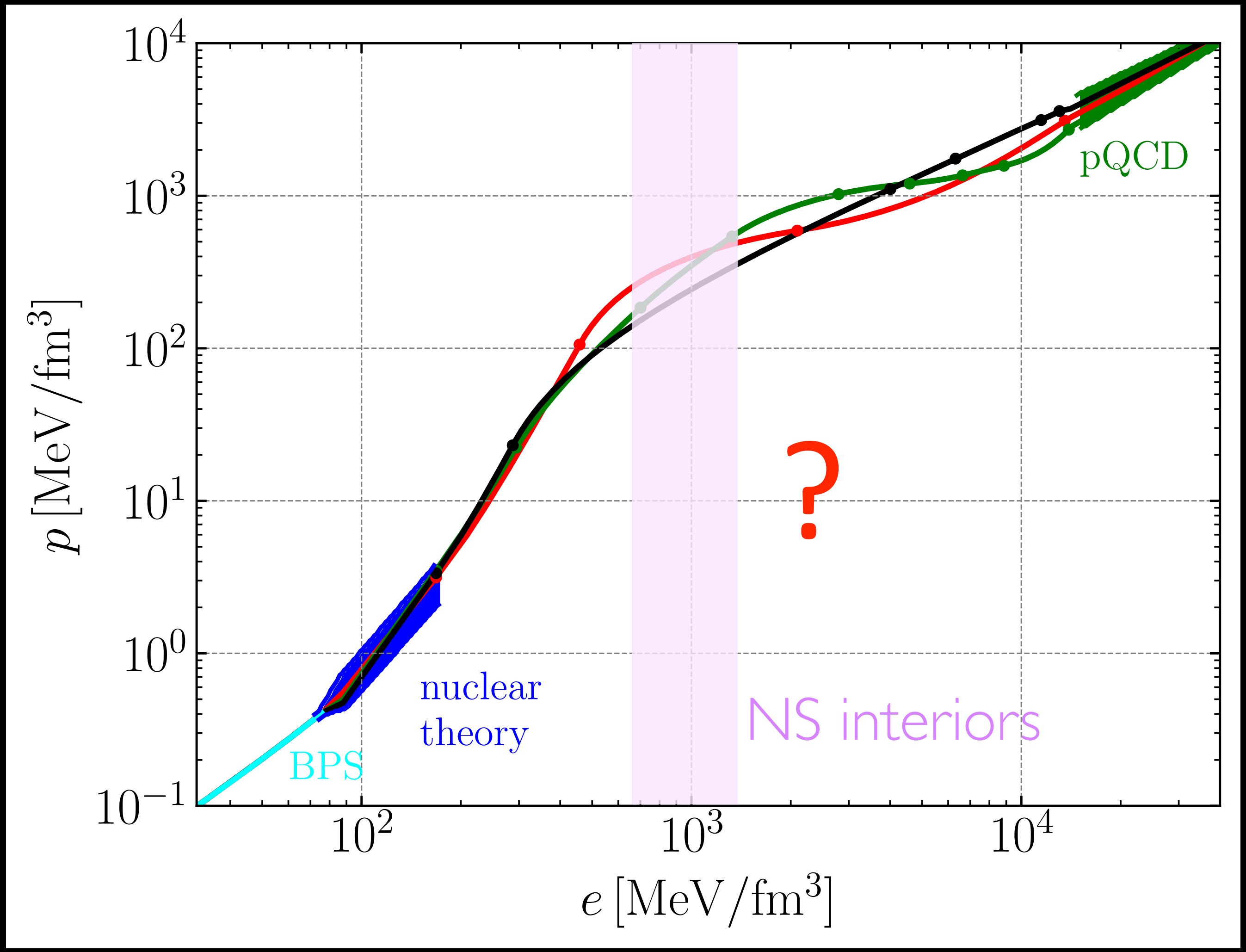
On the sound speed in neutron stars

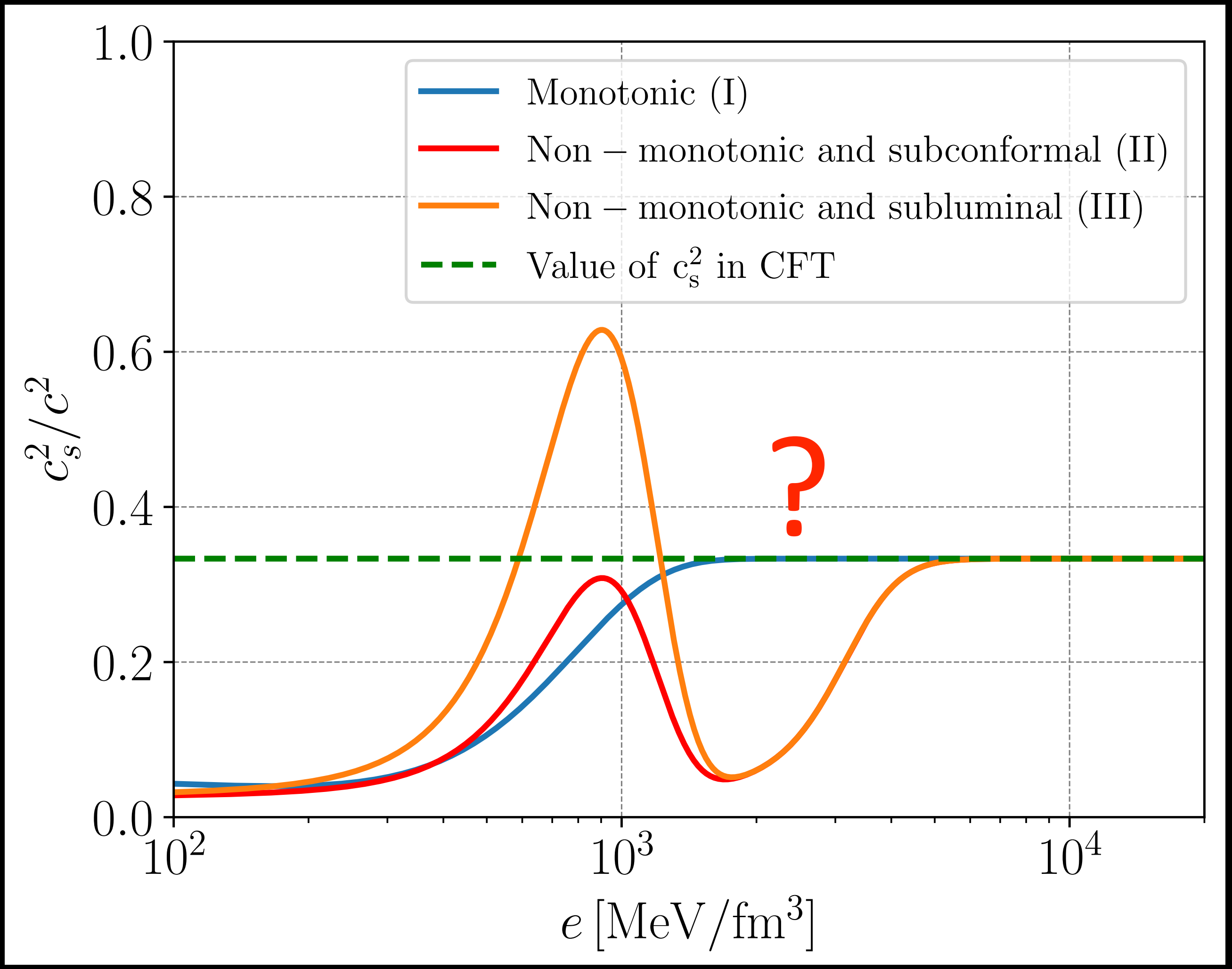


Altiparmak, Ecker, LR (2022)

A very basic question

The EOS of nuclear matter still remains an open question.
Some information is available but freedom is still large





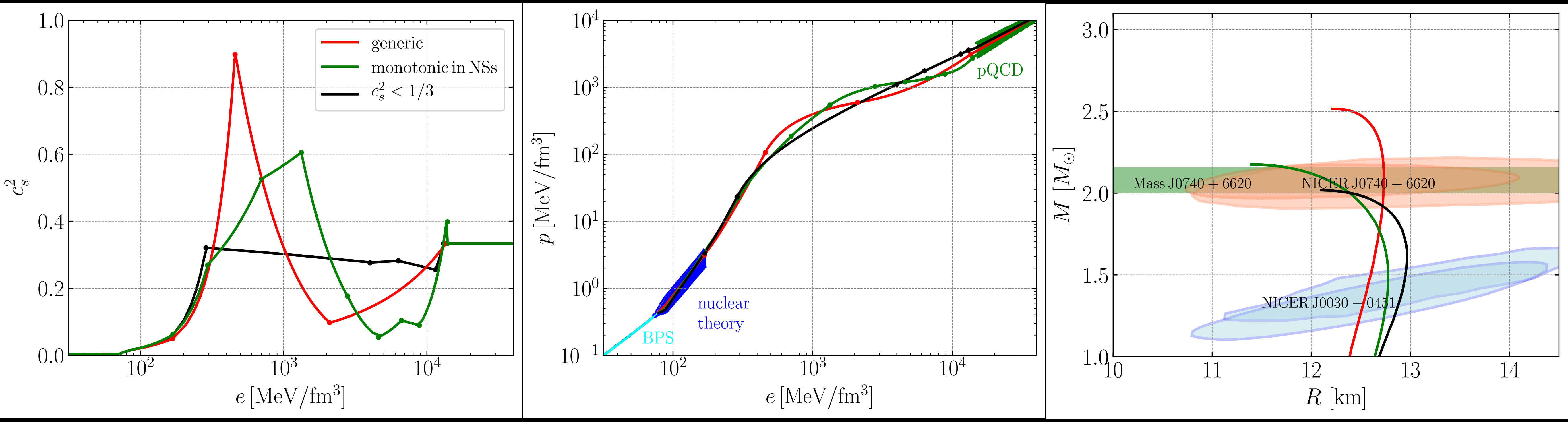
i) monotonic and sub-conformal: $c_s^2 < 1/3\,;$

ii) nonmonotonic and sub-conformal: $c_s^2 < 1/3;$

iii) nonmonotonic and sub-luminal: $c_s^2 < 1$

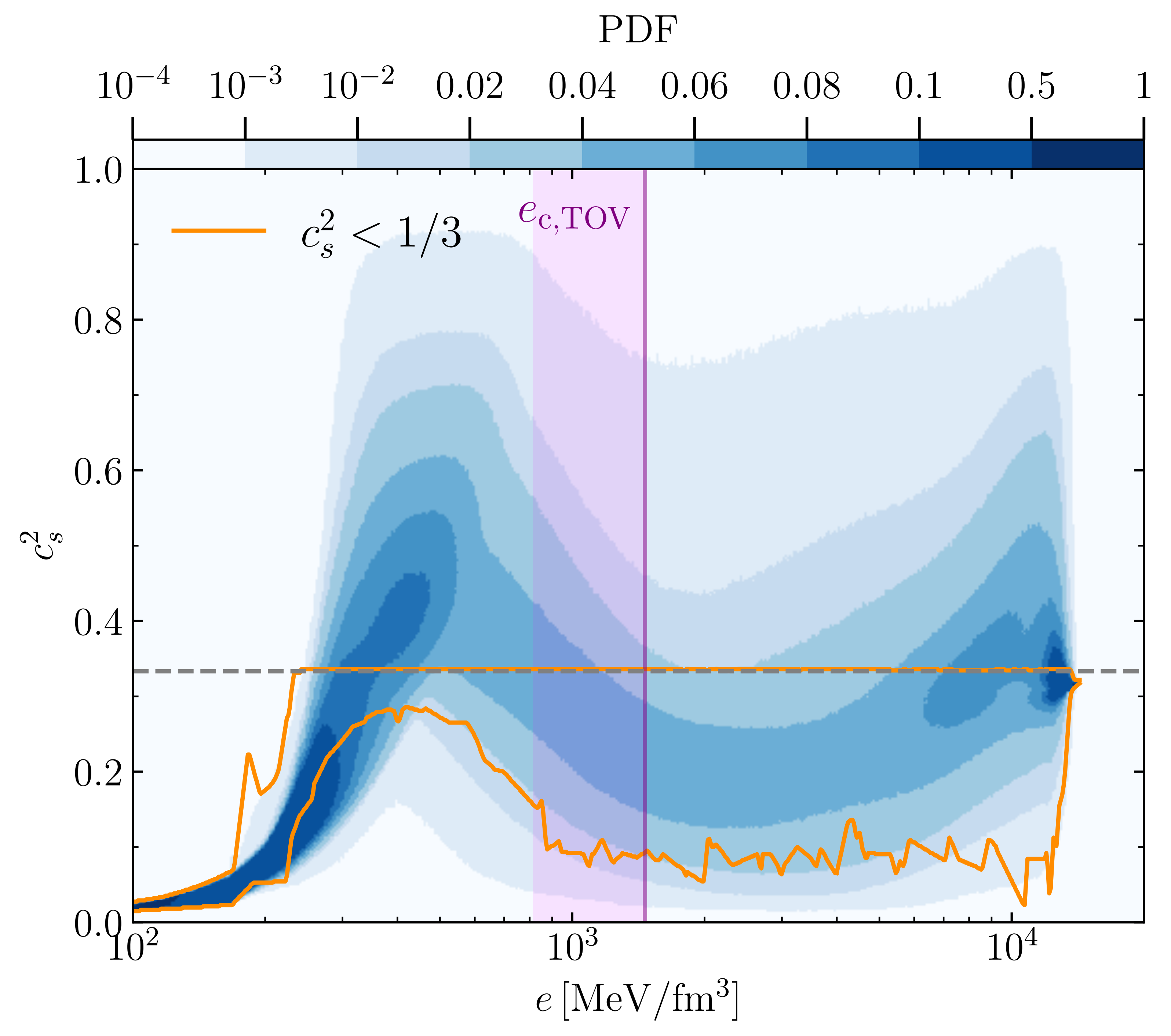
- Lacking stronger constraints, an agnostic approach is viable and followed by many (eg piecewise polytropes, Most+ 2018)

- Here, instead, we build an EOS starting from a piecewise prescription of the sound speed (7 segments are sufficient)



- Once an EOS is produced, we check it satisfies astrophysical constraints (max. mass, NICER limits). We repeat 1.5x10^7 times…
- In this way, ~ 10% of our EOSs survives and provides robust statistics from which we compute PDFs.

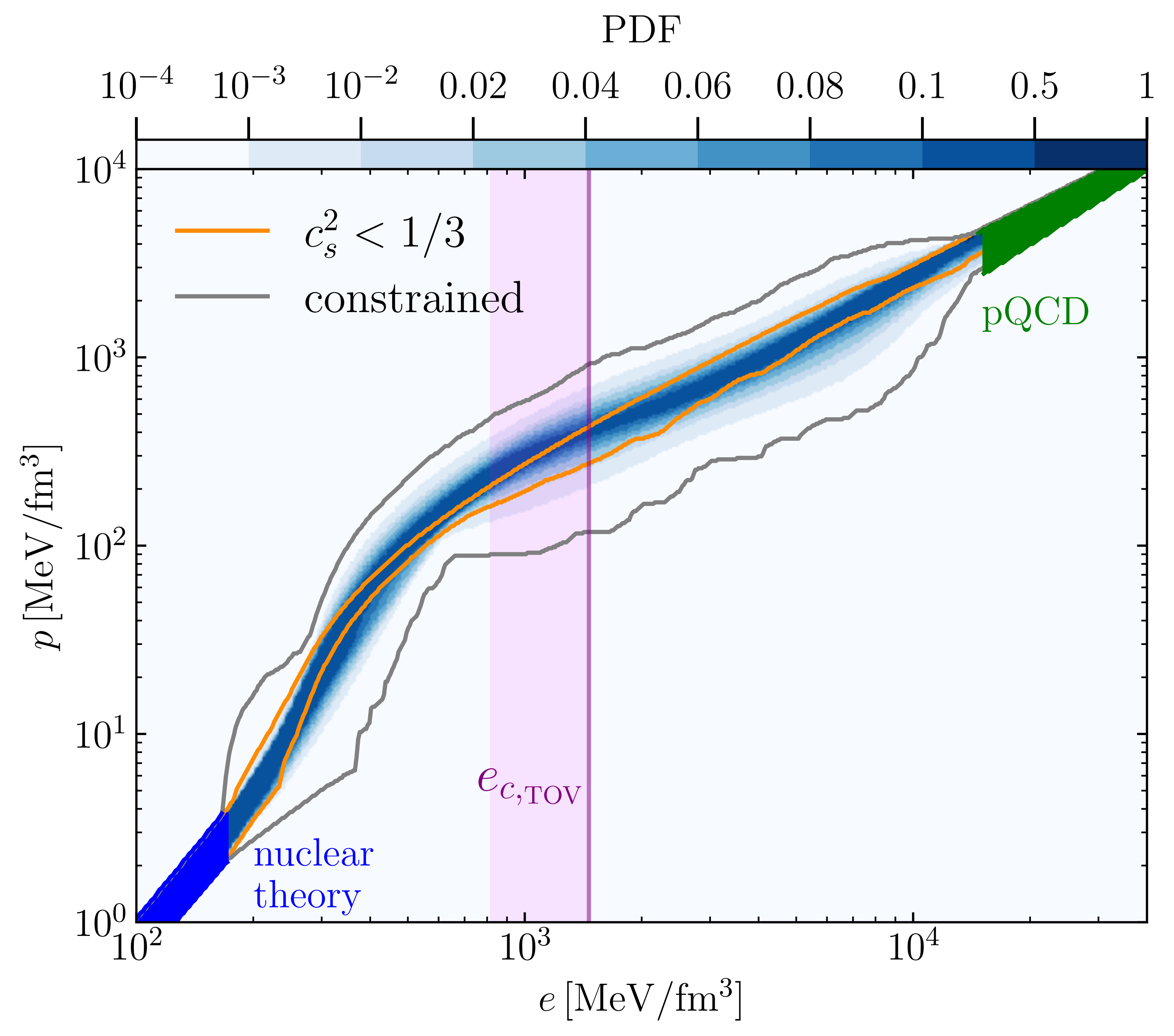
Sound speed PDF



Orange line marks region of sub-conformal EOSs (0.03%).

No monotonic sub-conformal EOS found.

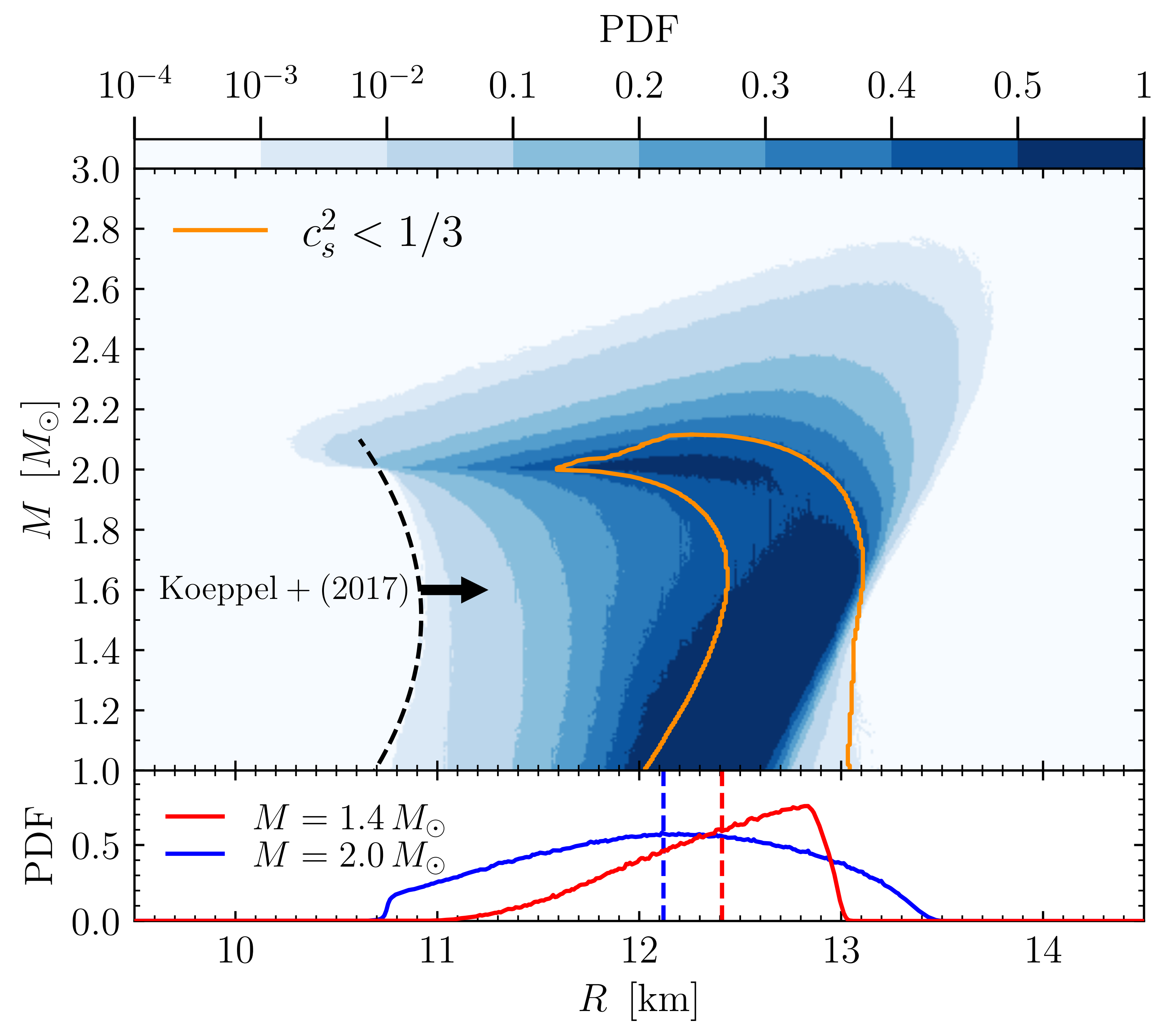
EOS PDF



Orange line marks region of sub-conformal EOSs (0.03%).

Note that 99% confidence region is very thin.

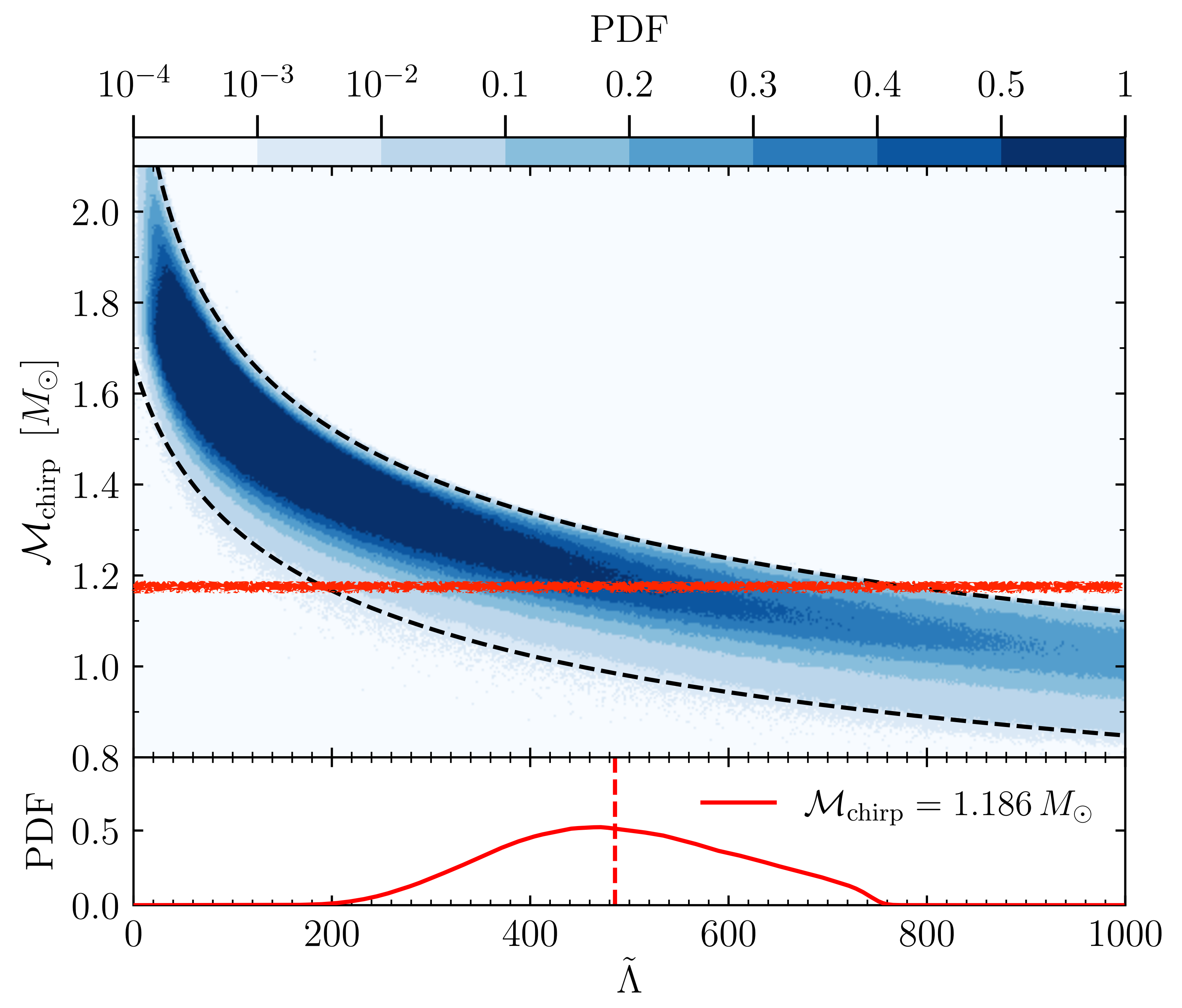
A more comprehensive picture



M-const. sections: $R_{1.4} = 12.42^{+0.52}_{-0.99}\,\mathrm{km}; \quad R_{2.0} = 12.12^{+1.11}_{-1.23}\,\mathrm{km}$

Lower bound on radii matches Köppel+ prediction from threshold mass.

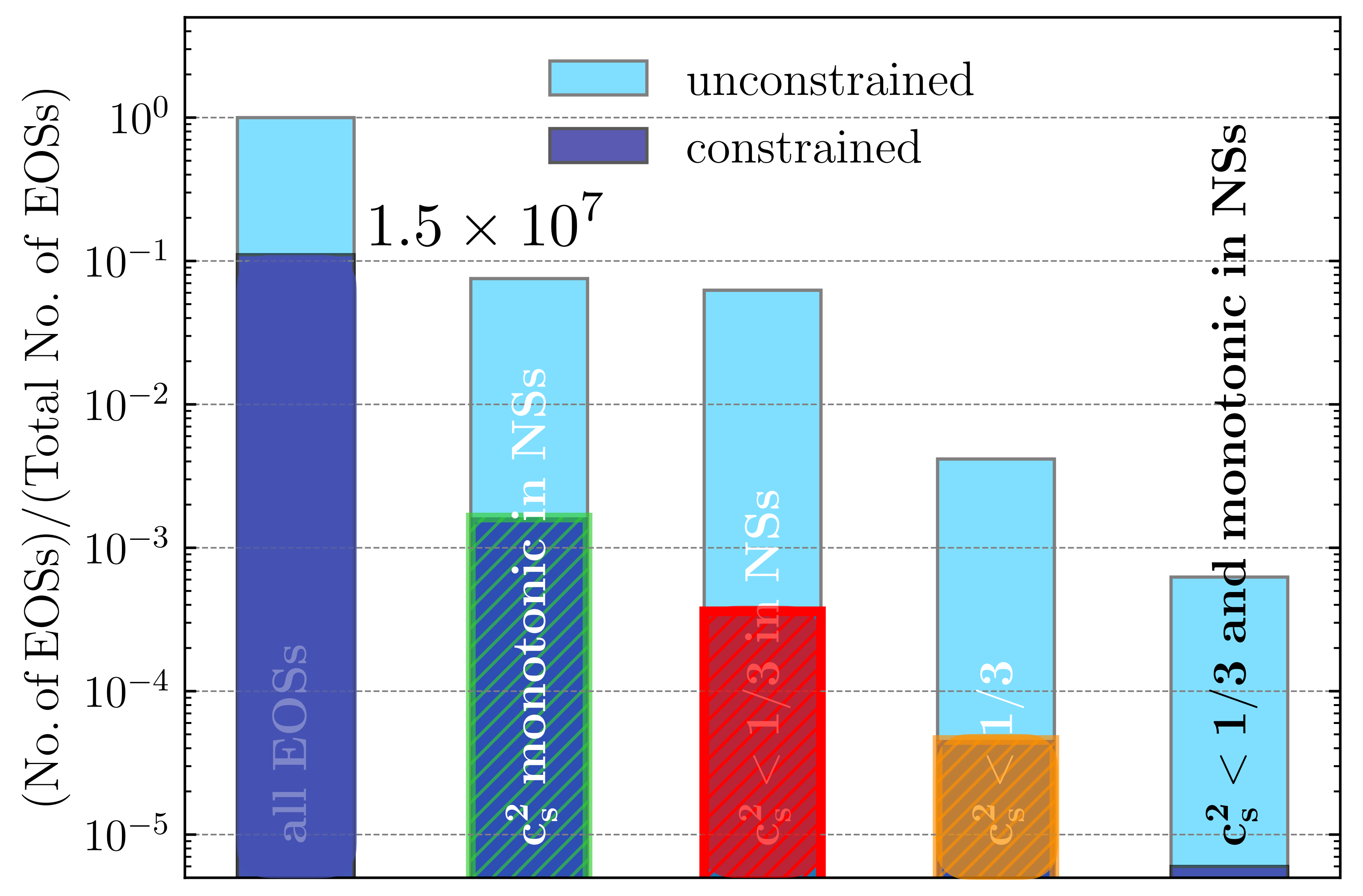
A more comprehensive picture



Simple behaviour of binary tidal deformability: $\tilde{\Lambda}_{\rm min\,(max)} = a + b\,\mathcal{M}_{\rm chirp}^{c}$

Straightforward bounds once a detection is made.

In summary



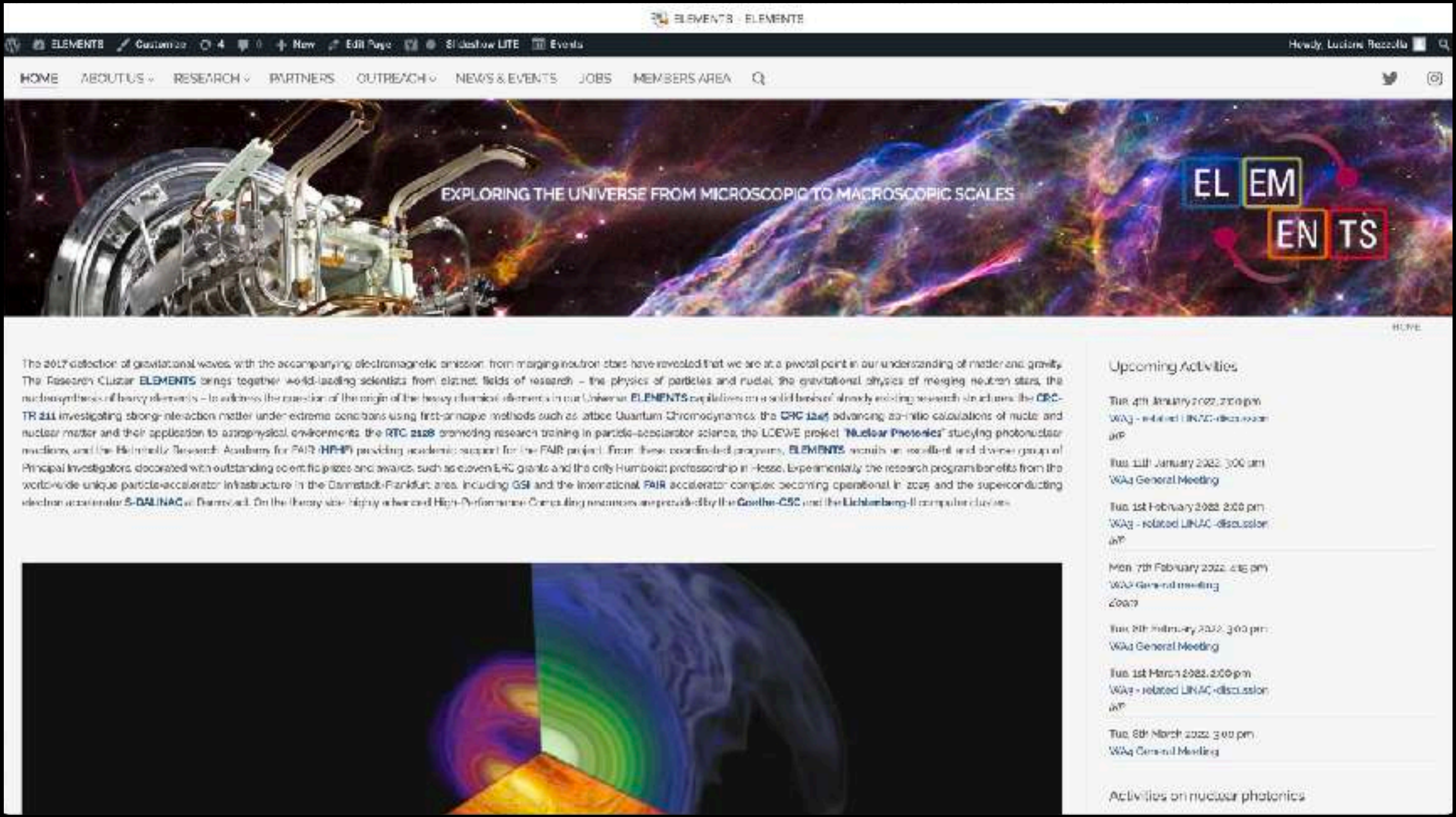
i) monotonic and sub-conformal: $c_s^2 < 1/3$; [0.004%]

ii) non-monotonic and sub-conformal in NSs: $c_s^2 < 1/3$; [0.03%]

iii) nonmonotonic and sub-luminal: $c_s^2 < 1$; [10%]

Much of the research presented is is part of **ELEMENTS**, an Hessian Research Cluster with Frankfurt Darmstadt and Giessen

Visit our site at: https://elements.science



The 2017 detection of gravitational waves with the accompanying electromagnetic emission from merging neutron stars have revealed that we are at a pivotal point in our understanding of matter and gravity. The Research Cluster **ELEMENTS** brings together world-leading scientists from distinct fields of research – the physics of particles and nuclei, the gravitational physics of merging neutron stars, the nucleosynthesis of heavy elements – to address the question of the origin of the heavy chemical elements in our Universe. **ELEMENTS** capitalizes on a solid basis of already existing research structures: the **CRC-TR 211** investigating strong-interaction matter under extreme conditions using first-principle methods such as lattice Quantum Chromodynamics, the **CRC 1245** advancing ab-initio calculations of nuclei and nuclear matter and their application to astrophysical environments, the **RTG 2128** promoting research training in particle-accelerator science, the LOEWE project "**Nuclear Photonics**" studying photonuclear reactions, and the Helmholtz Research Academy for FAIR (**HFHF**) providing academic support for the FAIR project. From these coordinated programs, **ELEMENTS** recruits an excellent and diverse group of Principal Investigators, decorated with outstanding scientific prizes and awards, such as eleven ERC grants and the only Humboldt professorship in Hesse. Experimentally, the research program benefits from the worldwide unique particle-accelerator infrastructure in the Darmstadt-Frankfurt area, including **GSI** and the international **FAIR** accelerator complex becoming operational in 2025 and the superconducting electron accelerator **S-DALINAC** at Darmstadt. On the theory side highly advanced High-Performance Computing resources are provided by the **Goethe-CSC** and the **Lichtenberg-II** computer clusters.

Upcoming Activities

Tue, 4th January 2022, 2:00 pm
WA3 - related LINAC-discussion
IKP

Tue, 11th January 2022, 3:00 pm
WA4 General Meeting

Tue, 1st February 2022, 2:00 pm
WA3 - related LINAC-discussion
IKP

Mon, 7th February 2022, 4:15 pm
WA4 General meeting
Zoom

Tue, 8th February 2022, 3:00 pm
WA4 General Meeting

Tue, 1st March 2022, 2:00 pm
WA3 - related LINAC-discussion
IKP

Tue, 8th March 2022, 3:00 pm
WA4 General Meeting

Activities on nuclear photonics

Conclusions

✳Spectra of post-merger shows peaks, some ”**quasi-universal**”.

✳When used together with tens of observations, they will set tight constraints on EOS: radius known with **~1 km** precision.

✳**GW170817** has already provided new limits on

$$2.01^{+0.04}_{-0.04} \le M_{\rm TOV}/M_\odot \le 2.16^{+0.17}_{-0.15}$$ **maximum mass**

$$12.00 < R_{1.4}/{\rm km} < 13.45 \quad \tilde{\Lambda}_{1.4} > 375$$ **radius, tidal deformability**

✳A **phase transition** after a BNS merger leaves GW **signatures** and opens a gate to access quark matter beyond accelerators.

✳**Sound speed** in neutron stars cannot be sub-conformal and monotonic; likely to be super-conformal somewhere in the interior.